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# Molecular outflows in local galaxies: Method comparison and a role of intermittent AGN driving

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## ABSTRACT

We report new detections and limits from a NOEMA and ALMA CO(1-0) search for molecular outflows in 13 local galaxies with high far-infrared surface brightness, and combine these with local universe CO outflow results from the literature. The CO line ratios and spatial outflow structure of our targets provide some constraints on the conversion steps from observables to physical quantities such as molecular mass outflow rates. Where available, ratios between outflow emission in higher J CO transitions and in CO(1-0) are typically consistent with excitation  $R_{11} \lesssim 1$ . However, for IRAS 13120–5453,  $R_{31} = 2.10 \pm 0.29$  indicates optically thin CO in the outflow. Like much of the outflow literature, we use  $\alpha_{\text{CO}(1-0)} = 0.8$ , and we present arguments for using  $C = 1$  in deriving molecular mass outflow rates  $\dot{M}_{\text{out}} = C \frac{M_{\text{out}} v_{\text{out}}}{R_{\text{out}}}$ . We compare the two main methods for molecular outflow detection: CO millimeter interferometry and *Herschel* OH-based spectroscopic outflow searches. For 26 sources studied with both methods, we find an 80% agreement in detecting  $v_{\text{out}} \gtrsim 150 \text{ km s}^{-1}$  outflows, and non-matches can be plausibly ascribed to outflow geometry and signal-to-noise ratio. For a published sample of 12 bright ultraluminous infrared galaxies with detailed OH-based outflow modeling, CO outflows are detected in all but one. Outflow masses, velocities, and sizes for these 11 sources agree well between the two methods, and modest remaining differences may relate to the different but overlapping regions sampled by CO emission and OH absorption. Outflow properties correlate better with active galactic nucleus (AGN) luminosity and with bolometric luminosity than with far-infrared surface brightness. The most massive outflows are found for systems with current AGN activity, but significant outflows in nonAGN systems must relate to star formation or to AGN activity in the recent past. We report scaling relations for the increase of outflow mass, rate, momentum rate, and kinetic power with bolometric luminosity. Short flow times of  $\sim 10^6 \text{ yr}$  and some sources with resolved multiple outflow episodes support a role of intermittent driving, likely by AGNs.

**Key words.** galaxies: kinematics and dynamics – galaxies: active – galaxies: ISM

## 1. Introduction

Gas outflows triggered by intense star formation and/or active galactic nuclei (AGNs) are important agents in the evolution of galaxies, in regulating the fraction of baryons inside a dark matter halo that are converted to stars, and for the metal enrichment of the circumgalactic and intergalactic medium (see, e.g.,

reviews by [Veilleux et al. 2005](#); [Fabian 2012](#)). A significant step forward was recently made in characterizing the molecular phase of outflows from local galaxies via detection of P-Cygni profiles of far-infrared OH lines (e.g., [Fischer et al. 2010](#); [Sturm et al. 2011](#); [Veilleux et al. 2013](#); [Spoon et al. 2013](#); [González-Alfonso et al. 2017](#)) and detection of broad line wings to the millimeter (mm) emission lines of CO (e.g.,

Feruglio et al. 2010; Alatalo et al. 2011; Cicone et al. 2014; Pereira-Santaella et al. 2018; Fluetsch et al. 2019, and references therein). Molecular outflows have been detected preferentially but not exclusively in dusty (ultra)luminous infrared galaxies ((U)LIRGs). Outflow rates  $\dot{M}_{\text{out}}$  up to several  $100 M_{\odot} \text{ yr}^{-1}$  and mass loading factors  $\dot{M}_{\text{out}}/\text{SFR}$  around and significantly above 1 have been observed. In extreme cases and if velocities are high enough to escape the galaxy, extrapolation of the current outflow rate would clear out most of the gas content of a galaxy within  $10^7 \text{ yr}$ . While both star formation and AGNs may contribute to driving outflows, the OH outflow velocity in ULIRGs as well as the CO-based outflow rates seem to be best correlated with AGN luminosity (Veilleux et al. 2013; Cicone et al. 2014; Fiore et al. 2017; Fluetsch et al. 2019). Ionized, atomic, and molecular gas phases may all significantly contribute to the outflow rates (Contursi et al. 2013; Rupke & Veilleux 2013; Janssen et al. 2016; Rupke et al. 2017), in this paper we focus on the usually dominant molecular phase. There is a continuum of properties from the fast and energetic outflows reported in the studies mentioned above, towards smaller-scale and typically slower and less energetic molecular outflows. These are increasingly reported from high spatial- and spectral-resolution ALMA and NOEMA interferometric studies of nearby galaxies (e.g., Combes et al. 2014, 2019; Querejeta et al. 2016; Salak et al. 2016; Slater et al. 2019; Ramakrishnan et al. 2019; Alonso-Herrero et al. 2019), along with indications of inflow, warps, and noncircular motions in complex, for example barred, potentials. Here, we focus on fast and energetic outflows with potential nonlocal impact.

Thanks to efficient *Herschel* spectroscopy, the most statistically robust studies currently are those concerning searches for OH P-Cygni absorptions. For those, the step from measuring outflow kinematics (e.g., Veilleux et al. 2013; Stone et al. 2016) to deriving outflow masses and outflow rates involves complex modeling (e.g., Sturm et al. 2011; González-Alfonso et al. 2017; Stone et al. 2018). For the CO-based searches, the obvious and highly desirable next steps are to extend the local galaxy samples under study, and to broaden the range of galaxy parameters covered. Additional value can come from truly spatially resolved studies of very nearby outflow sources.

Three major sources of uncertainty currently affect the derivation of molecular outflow rates from interferometric CO data: (i) the emission from outflowing gas has to be separated from that of the host; (ii) a CO conversion factor has to be adopted to convert from luminosity to gas mass, but may not necessarily be the same as for typical Galactic molecular clouds; and (iii) if only an outflow velocity and overall size are known, the conversion from outflow mass to outflow rate is uncertain by at least a factor of three due to different possible geometrical assumptions (see discussions for example in Cicone et al. 2014; Veilleux et al. 2017, and Sect. 4.2). Poorly known inclination of the outflow can add to these geometrical uncertainties.

Even if separated reliably from gas orbiting in the host galaxy, the outflowing gas will face a different fate depending on its velocity and the depth of the potential. There is a continuum from “fountains” that will re-join the host in much less than a gigayear to gas staying in the circumgalactic medium for longer periods until being re-accreted, and to gas able to escape the galaxy’s dark matter halo. In Sect. 3 below we conservatively identify outflowing gas with line wings only. This emphasizes the second and third paths, but at typical escape velocities of the order of  $600 \text{ km s}^{-1}$  the fraction of gas truly escaping a galaxy’s potential is typically small if ballistic motion is assumed (e.g., Fluetsch et al. 2019). Focus on fast wings outside the line core of

the galaxy also avoids ambiguities between more modest velocity outflow and other noncircular motions such as inflow or flows in the potential of a stellar bar.

Using *Herschel*  $70 \mu\text{m}$  data, Lutz et al. (2016, 2018) derive sizes of the far-infrared emission region for large samples of local galaxies, in order to derive scaling relations of far-infrared size and surface brightness with salient galaxy properties, and establish that AGN hosts preferentially have more compact star formation than inactive galaxies. This sample can be used for an objective selection of compact far-infrared sources (Table 1 in Lutz et al. 2016), based on observations in the *Herschel* archive that are not affected by bright neighbors. The 18 sources listed in Table 1 of Lutz et al. (2016) include several for which CO outflows have been recently detected, as one might expect for their high far-infrared surface brightness  $\Sigma_{\text{FIR}} \gtrsim 10^{11.75} L_{\odot} \text{ kpc}^{-2}$ . If star formation powers the far-infrared, this is well beyond thresholds in star formation rate surface density above which star formation-driven outflows from galaxy disks tend to be observed in the local universe (Heckman 2002) or at high redshift (Newman et al. 2012; Davies et al. 2019). Alternatively, if an AGN significantly contributes to the observed far-infrared emission, it must be luminous and could drive an outflow. The likelihood of hosting outflows, and the proximity of the sample ( $z = 0.0072\text{--}0.0676$  with median 0.027) make it an obvious basis for a search for local outflow laboratories, well suited for detailed study.

We here report on a NOEMA and ALMA search for CO(1-0) outflows in galaxies from the Lutz et al. (2016) sample of compact far-infrared galaxies that lack previous deep CO data, and discuss the results in conjunction with literature results for the remainder of the sample and outside the sample’s definition criteria. We adopt an  $\Omega_{\text{m}} = 0.3$ ,  $\Omega_{\Lambda} = 0.7$  and  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$  cosmology, redshift-independent distances from the NASA Extragalactic Database (NED), if available, for  $z < 0.01$  galaxies, a Chabrier (2003) initial mass function (IMF), a conversion for far-infrared luminosity to star formation rate (SFR)  $\text{SFR} = 1.9 \times 10^{-10} L_{\text{FIR}} = 40\text{--}120 \mu\text{m}$  as appropriate for the Kennicutt (1998) conversion scaled to Chabrier IMF, and  $L_{\text{IR}}(8\text{--}1000 \mu\text{m})/L_{\text{FIR}}(40\text{--}120 \mu\text{m}) = 1.9$ .

## 2. Sample and observations

### 2.1. Northern Extended Millimeter Array observations

Observations for the  $\text{Dec} > 10^\circ$  targets were taken with the IRAM Northern Extended Millimeter Array (NOEMA) on Plateau de Bure between May 2016 and February 2018 (Projects s16bh, s17ar, w17cp). The CO(1-0) line was observed in the 3 mm band, mostly using the WideX correlator (3.6 GHz bandwidth) and typically 6 or 7 antennae in configurations between C and D. These low-spatial-resolution ( $2\text{--}4''$ ) configurations were chosen to initially provide maximum detection sensitivity, and not over-resolve faint and possibly extended high-velocity outflow components.

Table 1 lists the baselines used, achieved RMS sensitivity per beam for a  $20 \text{ km s}^{-1}$  channel, and beam size. For III Zw 035, a strong outflow was detected and already spatially resolved in the low-resolution data. This source was followed up in the higher-resolution A configuration using the new PolyFix correlator and new receivers. The combined data yield the smaller beam listed for this source in Table 1.

Calibration and mapping followed standard procedures in the GILDAS<sup>1</sup> environment (Guilloteau & Lucas 2000). For

<sup>1</sup> <http://www.iram.fr/IRAMFR/GILDAS>

**Table 1.** NOEMA and ALMA CO(1-0) observations.

Source	$cz_h$ (km s <sup>-1</sup> )	RA J2000	Dec J2000	Date(s)	Min, max baseline (m)	rms (20 km s <sup>-1</sup> ) (mJy beam <sup>-1</sup> )	Beam (arcsec)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
III Zw 035	8221	01:44:30.71	17:06:07.6	07/2016–02/2018	18.4, 760.0	0.30	$1.42 \times 0.94$
NGC 2623	5524	08:38:24.08	25:45:16.7	05/2017–09/2017	15.0, 174.1	0.75	$4.03 \times 3.33$
IC 860	3889	13:15:03.53	24:37:07.9	08/2016–11/2016	15.0, 303.4	0.92	$2.81 \times 1.79$
NGC 7479	2387	23:04:56.65	12:19:22.3	05/2016–06/2016	15.0, 262.0	1.15	$2.98 \times 2.27$
IRAS F05189–2524	12843	05:21:01.47	–25:21:45.4	09/2015	15.1, 1600	0.37	$0.52 \times 0.47$
IRAS F09111–1007 W	16276	09:13:36.38	–10:19:28.2	12/2016	15.1, 460.0	0.33	$1.84 \times 1.47$
IRAS F10173+0828	14687	10:20:00.26	08:13:33.6	12/2016	15.1, 460.0	0.34	$1.84 \times 1.68$
IRAS F12224–0624	7938	12:25:03.91	–06:40:52.0	12/2016	15.1, 491.0	0.37	$1.76 \times 1.53$
NGC 4418	2099	12:26:54.65	–00:52:38.4	01/2017	15.1, 460.0	0.58	$1.62 \times 1.19$
IRAS 13120–5453	9308	13:15:06.20	–55:09:24.0	07/2015–08/2016	15.1, 1800	0.65	$0.54 \times 0.52$
IRAS F14378–3651	20417	14:40:59.01	–37:04:32.0	07/2015	43.3, 1600	0.67	$0.57 \times 0.51$
IRAS F17207–0014	12825	17:23:21.98	–00:17:01.7	01/2017	15.1, 452.8	0.46	$2.52 \times 1.87$
IRAS F20551–4250	12891	20:58:26.79	–42:39:00.3	06/2015	43.3, 1600	0.64	$0.57 \times 0.45$

**Notes.** Observations obtained with NOEMA (top) and ALMA (bottom). (1) Source name. (2) Heliocentric radial velocity based on our CO(1-0) data. (3), (4) Coordinates of adopted phase center. See Table 2 for the measured position of mm continuum. (5) Observation period. (6) Interferometer baselines used. (7) rms sensitivity for a 20 km s<sup>-1</sup> channel width. (8) Beam size.

**Table 2.** Continuum properties.

Source	RA J2000	Dec J2000	E(RA) (arcsec)	E(Dec) (arcsec)	<i>FWHM</i> (arcsec)	$\nu_{\text{Obs}}$ (GHz)	$S_{\nu, \text{Obs}}$ (mJy)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
III Zw 035	01:44:30.538	17:06:08.68	0.01	0.01	$0.43 \pm 0.01$	112.2	$3.84 \pm 0.03$
IRAS F05189–2524	05:21:01.401	–25:21:45.28	0.01	0.01	$0.17 \pm 0.03$	104.1	$2.40 \pm 0.03$
NGC 2623	08:38:24.087	25:45:16.67	0.01	0.02	$1.22 \pm 0.06$	113.2	$6.40 \pm 0.08$
IRAS F09111–1007 W	09:13:36.452	–10:19:30.00	0.01	0.01	$0.76 \pm 0.11$	102.3	$1.33 \pm 0.03$
IRAS F09111–1007 E	09:13:38.829	–10:19:19.98	0.01	0.01	$<3.12$	102.3	$0.62 \pm 0.12$
IRAS F10173+0828	10:20:00.203	08:13:33.93	0.01	0.01	$<0.75$	102.9	$1.72 \pm 0.03$
IRAS F12224–0624	12:25:03.909	–06:40:52.63	0.01	0.01	$<0.48$	105.8	$2.33 \pm 0.33$
NGC 4418	12:26:54.610	–00:52:39.45	0.01	0.01	$0.17 \pm 0.05$	107.7	$11.13 \pm 0.04$
IC 860	13:15:03.505	24:37:07.70	0.01	0.01	$<0.54$	113.8	$6.64 \pm 0.08$
IRAS 13120–5453	13:15:06.328	–55:09:22.80	0.01	0.01	$0.89 \pm 0.03$	105.3	$10.07 \pm 0.16$
IRAS F14378–3651	14:40:59.013	–37:04:31.91	0.01	0.01	$0.37 \pm 0.05$	101.4	$1.47 \pm 0.06$
IRAS F17207–0014	17:23:21.959	–00:17:00.91	0.01	0.01	$0.66 \pm 0.06$	104.1	$8.70 \pm 0.06$
IRAS F20551–4250	20:58:26.792	–42:39:00.29	0.01	0.01	$0.25 \pm 0.03$	104.1	$2.78 \pm 0.05$
NGC 7469	23:04:56.634	12:19:22.67	0.04	0.04	$0.89 \pm 0.15$	114.4	$1.77 \pm 0.09$

**Notes.** Results are from UV fitting a Gaussian model to the NOEMA continuum visibilities, and from fitting a Gaussian model to the ALMA continuum image. IRAS F09111–1007E is not part of our sample but detected along with IRAS F09111–1007W. (1) Source name. (2), (3) Continuum position. (4), (5) Uncertainties of position. These include only the formal fit errors, to which 2–5% of the beam size should be added to account for phase calibration stability (Warmels et al. 2018). (6) Intrinsic source full width at half maximum FWHM (beam-deconvolved). (7) Observed frequency. (8) Total continuum flux density from Gaussian fit.

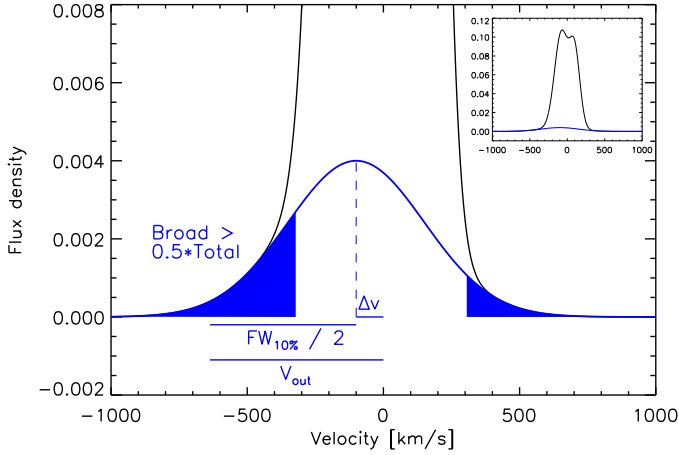
CO-bright and compact sources (III Zw 035, IC 860, NGC 2623), self calibration was able to improve the dynamic range in the channels with bright line emission, but this has little impact on the detection of outflow wings and continuum. Velocities quoted below are relative to our CO-based heliocentric redshifts as listed in Table 1. Millimeter continuum emission was clearly detected for all sources. Positional offsets quoted below are relative to the continuum positions that are listed along with other continuum properties in Table 2. The continuum was subtracted from the visibilities, and all line information is based on continuum free visibilities or spectral cubes. The observations were targeted at CO(1-0), for which we derived spectral cubes at 20,

60, and 200 km s<sup>-1</sup> resolution; occasionally we also used data for velocity ranges specifically tailored to the outflow properties of a given source. Total fluxes for CO and some other transitions detected in the wide spectral bandwidth are listed in Table B.1.

## 2.2. Atacama Large Millimeter/submillimeter Array observations

The Dec < 10° targets were observed with the Atacama Large Millimeter/submillimeter Array (ALMA) Cycle 4 program 2016.1.00177.S in late 2016 to early 2017, using between





**Fig. 1.** Visualization of adopted conventions for outflow velocity and outflow CO flux.  $v_{\text{out}} = |\Delta v| + FW_{10\%}/2$  (Eq. (1)) is derived from the shift of the broad Gaussian outflow component with respect to systemic velocity and its full width at a tenth of its peak.  $S_{\text{CO}}$  (gauss) is the integral of this broad component. The flux in the wings of the broad component  $S_{\text{CO}}$  (wings), indicated by the blue shaded area, integrates the broad component only over the velocity ranges for which it contributes at least 50% of the total flux density of the line.

38 and 46 antennae. Again, the CO(1-0) line was observed in a relatively compact configuration yielding an  $\approx 2''$  beam in ALMA Band 3, and observational parameters are listed in Table 1. The pipeline-calibrated data were mapped using standard CASA (McMullin et al. 2007) procedures. Three more sources from Table 1 of Lutz et al. (2016) were already observed in CO(1-0) in ALMA program 2012.1.00306.S (PI E. Sturm) using  $\sim 0.5''$  beams. For IRAS 13120–5453, we combined our 2012.1.00306.S CO(1-0) data with a similar dataset from 2015.1.00287.S (PI K. Sliwa), in order to improve signal-to-noise ratio. Again, we mapped the calibrated data in CASA and include the fourth target of 2012.1.00306.S, IRAS F20511–4250, which falls just below the criteria of Table 1 of Lutz et al. (2016).

In reducing the ALMA data, we adopted the same channel widths as for the NOEMA data. Self calibration was applied for IRAS F09111–1007W, NGC 4418, IRAS F17207–0014, IRAS F20551–4250 but not the other sources for which we found no noticeable improvement.

### 3. Outflow properties of the targets

First, we determine whether an outflow is present, and then measure its CO flux, velocity, and size. Outflow identification can be relatively straightforward in cases where an extremely broad outflow line component is superposed on a core of velocity width that is typical for undisturbed gas moving in the gravitational potential of a galaxy. This distinction is more difficult if the velocity range with observed “broad” CO emission does not strongly exceed plausible nonoutflow rotational velocities within the galaxy’s potential. Different radial distributions of molecular gas disks rotating in the potential, misaligned disk components, bar flows, or gas moving in the potential of a merger all can lead to relatively complex total line profiles over a velocity range that is similar to local circular velocities. For outflows from spatially resolved disks in regular rotation, the kinematic separation of outflowing and nonoutflowing gas can be improved by removing the regular velocity field (e.g., Newman et al. 2012), but this approach is less suited for a

sample of compact nuclei that also includes mergers. Needing two Gaussian components for a satisfactory fit of the line profile is alone not sufficient to ascribe the broader component to an outflow.

We hence continue with a brief discussion of individual targets on the basis of total CO maps and selected channel maps, spectra, positions in different velocity channels, and position-velocity diagrams, and visualize in Figs. A.1–A.13 such basic CO(1-0) properties. As appropriate for the observed spatial extent of an outflow, we extract nuclear or aperture spectra and fit multiple Gaussians to the spectrum, which is the sum of host emission and (if present) broad outflow. Visual inspection is used to choose the number of Gaussians needed to fit the host emission. Figure 1 illustrates the adopted conventions for outflow velocity and flux. For intrinsically somewhat asymmetric outflows and little obscuration, the shift  $\Delta v$  of the broad component can take positive as well as negative values, and this is indeed observed (Table 3). Unless mentioned otherwise, we adopt as outflow velocity

$$v_{\text{out}} = |\Delta v| + \frac{FW_{10\%}}{2}, \quad (1)$$

which is computed from the shift of the centroid of the broad Gaussian line component for the outflow with respect to the systemic velocity, and the full width at a tenth of the maximum of this broad outflow component. With this definition,  $v_{\text{out}}$  will be in the line profile wings that are dominated by the outflow. Because outflow components at velocities that also show host CO emission are hard to quantify, we quote below in Table 3 two measured CO fluxes for the outflow:  $S_{\text{CO}}$  (gauss) is the integral of the full Gaussian broad line component, while  $S_{\text{CO}}$  (wings) integrates this Gaussian only over the range(s) where the outflow dominates emission (i.e., contributes more than 50% of the total fitted profile). In further analysis, we adopt the more conservative  $S_{\text{CO}}$  (wings). This is similar to the integration of wings only (e.g., Ciccone et al. 2014) or to removal of host flux on the basis of a disk model (e.g. Veilleux et al. 2017; Herrera-Camus et al. 2019), but different from adopting the full broad Gaussian flux as in several other references in the literature. For objects without outflow detection, we assign limits for the outflow CO flux:  $S_{\text{CO, wings}} < 10 \times \text{rms}_{200} \times 200$ , based on the rms noise ( $\text{Jy beam}^{-1}$ ) over a  $200 \text{ km s}^{-1}$  channel. Using a factor ten here rather than three considers that undetected outflows may be spatially extended or spectrally broad. The outflow radius is directly taken from the maps for a few well-resolved outflows, otherwise we derive a radius

$$R_{\text{out}} = |\Delta R| + FWHM/2 \quad (2)$$

Here,  $|\Delta R|$  is the spatial shift between line position and continuum position and FWHM refers to the line emission. Both line position and FWHM are derived from a Gaussian spatial fit in the UV plane using a velocity range that is dominated by outflow, that is, by the broad component of the spectral multi-Gaussian fit.  $R_{\text{out}}$  is effectively a beam-deconvolved intrinsic radius. For symmetric bipolar outflow we use average properties of the two lobes. We do not attempt to correct the measured  $R_{\text{out}}$  or  $v_{\text{out}}$ , or the literature results adopted in Sect. 3.2, for the effects of outflow opening angle and inclination. This emphasizes homogeneity of results with respect to more complete information that may be available for a few of the best spatially resolved and strongly collimated outflows. Observed outflow quantities or limits are listed in Table 3.

**Table 3.** Measured CO(1-0) outflow parameters.

Source	$S_{\text{CO}}$ (gauss) (Jy km s <sup>-1</sup> )	$S_{\text{CO}}$ (wings) (Jy km s <sup>-1</sup> )	$\Delta v$ (km s <sup>-1</sup> )	$FW_{10\%}$ (km s <sup>-1</sup> )	$v_{\text{out}}$ (km s <sup>-1</sup> )	$R_{\text{out}}$ ( $''$ )
III Zw 035	41.05	14.25	23	579	313	4.00
IRAS F05189–2524	2.85	1.11	60	771	446	1.24
NGC 2623	3.21	0.92	–33	1108	587	1.62
NGC 2623 fossil		0.37			400	14.00
IRAS F09111–1007 W	3.34	0.95	–9	644	331	0.58
IRAS F10173+0828		<0.22				
IRAS F12224–0624	0.84	0.35	2	1128	566	0.50
NGC 4418	2.53	0.43	9	630	324	0.23
IC 860		<0.58				
IRAS 13120–5453	34.50	4.29	1	588	301	0.63
IRAS F14378–3651		<0.43				
IRAS F17207–0014	0.84	0.49	–514	1203	1116	0.45
IRAS F20551–4250	0.98	0.40	96	787	490	0.27
NGC 7479		<0.73				

**Notes.** See Sect. 3 for definitions and methods to measure these quantities.

### 3.1. Properties of individual targets

**III Zw 035** (Fig. A.1). Extended CO emission is centered on the brighter northern lenticular-shaped component of this galaxy pair (for an optical image see, e.g., Kim et al. 2013). The second fainter optical component, offset roughly along the major axis of the brighter galaxy by  $\sim 7''$  toward PA  $\sim 200^\circ$  does not stand out in CO emission or kinematics. On top of this extended disk-like CO emission with a weak SW–NE velocity gradient that does not strictly follow the major axis of the galaxy, we detect a strong and asymmetric spatially resolved bipolar CO outflow up to velocities of  $\pm 500$ – $600$  km s<sup>-1</sup>. The outflow has PA  $\sim -45^\circ$  and we adopt as  $R_{\text{out}}$  the  $4''$  extent of its stronger NW (redshifted) side.

The line profile in an  $r = 3''$  aperture, offset by  $1''$  each to N and W, includes both sides of the outflow and is well fit by the sum of three Gaussians. One is narrow and the other two are broad and (on average) redshifted, with  $FWHM$   $270$  km s<sup>-1</sup> and  $633$  km s<sup>-1</sup>. We assign both broad components to the outflow. The sum of the two broad Gaussians dominates the flux within this aperture at all velocities. Hence, it is not possible to follow our standard procedure and separate the outflow line wings from the line core at a velocity where the outflow starts to dominate the line profile. Instead, guided by the PV diagrams, we adopt the velocity range from  $-120$  km s<sup>-1</sup> to  $120$  km s<sup>-1</sup> as line core and include into our “wings” CO-flux-only emission from both broad components that is outside that range. The adopted  $\Delta v$ ,  $FW_{10\%}$ , and  $v_{\text{out}}$  are the flux-weighted means of the respective values for the two broad components.

**IRAS F05189–2524** (Fig. A.2). Extended CO emission centered on the nucleus shows only little velocity change on  $1$ – $2''$  scales, consistent with a mostly face-on view of this Type 1 AGN host. A gradient from approaching (W) to receding (E) velocities is seen in the inner  $0.3''$ . Beyond the  $\sim \pm 200$  km s<sup>-1</sup> covered by these host structures, a clear red wing reaches out to  $\sim 500$  km s<sup>-1</sup> and traces an outflow. This emission is extended and we use an  $r = 1''$  aperture to extract its flux (green component in the multi-Gaussian fit shown top right in Fig. A.2). There may be a hint of near-nuclear blueshifted emission down to  $\sim -1000$  km s<sup>-1</sup>, but this never reaches  $>3\sigma$ , and hence is not discussed further. We adopt  $v_{\text{out}} = 446$  km s<sup>-1</sup> from the Gaussian fit to the  $r < 1''$  outflow spectrum, and  $R_{\text{out}} = 0.41'' + 0.5 \times 1.66''$  from a Gaussian spatial fit to the  $210$ – $510$  km s<sup>-1</sup> visibilities.

**NGC 2623** (Fig. A.3). The inner  $2''$  show a roughly E–W velocity gradient from gas rotating around the nucleus of this merger. At  $|v| > 300$  km s<sup>-1</sup>, line wings trace a bipolar outflow in approximately N–S (polar) direction. We adopt  $v_{\text{out}} = 587$  km s<sup>-1</sup> from the fit to the  $r < 4''$  outflow spectrum and  $R_{\text{out}} = 1.62''$  from averaging the properties of the blue and red outflows at  $|v| = [300, 500]$  km s<sup>-1</sup>.

In addition to the bipolar outflow with its blue component shifted to the N, the  $[-500, -300]$  km s<sup>-1</sup> channel shows emission at  $\sim 14''$  from the nucleus, at PA  $\sim 130^\circ$ . This emission peaks at  $4.2\sigma$  and has total flux  $0.37$  Jy km s<sup>-1</sup>. Given the relatively distinct morphology, we do not assign this emission to the main bipolar outflow but interpret it as the result of an earlier outflow episode with different orientation (see also Sect. 5.3). We list this outflow as “NGC 2623 fossil”.

**IRAS F09111–1007W** (Fig. A.4). The inner  $2''$  show a roughly E–W velocity gradient from gas rotating around the nucleus of this galaxy, which is part of a wide pair. At  $|v| \gtrsim 200$  km s<sup>-1</sup>, line wings trace a bipolar outflow in approximately N–S (polar) direction. We adopt  $v_{\text{out}} = 331$  km s<sup>-1</sup> from the fit to the nuclear outflow spectrum and  $R_{\text{out}} = 0.58''$  from averaging the properties of the blue and red outflows at  $|v| = [210, 510]$  km s<sup>-1</sup>.

**IRAS F10173+0828** (Fig. A.5). A rotating disk with kinematic major axis at PA  $\sim 45^\circ$  is seen out to a radius of  $\sim 8''$ . The bright and compact nuclear CO emission does not obviously connect to this rotating disk and has wings out to  $\pm 300$  km s<sup>-1</sup>, beyond the  $\pm 100$  km s<sup>-1</sup> LOS velocities of the disk. While this could indicate a modest velocity outflow, we cannot firmly exclude bound motion in the near-nuclear potential and hence assign an upper limit to the emission of a faster molecular outflow.

**IRAS F12224–0624** (Fig. A.6). Weak blue- and redshifted outflow components are  $4\sigma$  detected in addition to bright spatially compact CO that has a velocity gradient along PA  $\sim -30^\circ$ . We adopt  $v_{\text{out}} = 566$  km s<sup>-1</sup> from the fit to the nuclear outflow spectrum. The spatial width of the weak outflow is very poorly constrained, we adopt a fiducial  $R_{\text{out}} = 0.5''$ .

**NGC 4418** (Fig. A.7). CO emission shows a compact core that is on average redshifted by  $\approx 15$  km s<sup>-1</sup> relative to the extended emission. The redshift in Table 1 refers to extended

emission at radii of 3–8". At first glance, the nuclear spectrum and PV diagram suggest a strong red and a weak blue wing to CO(1-0). But the compact, warm, and dense nuclear region of NGC 4418 is known to have a rich chemistry and a complex mm or submm spectrum (e.g., Sakamoto et al. 2013; Costagliola et al. 2015), meaning that sensitive ALMA data detect rare molecules in addition to CO and CN, which are detected for all sources of the sample within the frequency range of our data. Specifically, there is emission near  $\sim -740 \text{ km s}^{-1}$  relative to CO, which we identify with emission from the nitrogen sulfide radical NS,  $J = 5/2-3/2$  at  $\sim 115.55 \text{ GHz}$  rest frequency (see also line surveys such as Aladro et al. 2015; Meier et al. 2015). This includes six hyperfine components, and  $\Lambda$ -doubling means that NS emission of a similar strength is superposed on the red wing of CO at  $\sim 290 \text{ km s}^{-1}$ , with another six hyperfine components at  $\sim 115.15 \text{ GHz}$  rest. We use the "blue" complex and local thermal equilibrium line ratios for the 6 + 6 lines in the blue and red complex (M. Drosdovskaya, priv. comm.) to infer and subtract contamination by NS for the red wing of CO. The adopted intensities are for  $T = 150 \text{ K}$  but the relative strength of blue and red complex varies by less than 1% between 100 and 200 K, the range of central dust temperatures suggested for NGC 4418 by Sakamoto et al. (2013). We fix relative fluxes of the NS line components and their velocities relative to CO, but vary the overall flux normalization and line width to fit the blue nitrogen sulfide complex, and then subtract both complexes from the spectrum (see Fig. A.7).

Indeed much of the wing redwards of the CO(1-0) line is contributed by nitrogen sulfide, but CO line wings out to  $\gtrsim \pm 300 \text{ km s}^{-1}$  remain on both the blue and red sides after subtraction of that contamination. We proceed with the NS-subtracted spectrum and adopt  $v_{\text{out}} = 324 \text{ km s}^{-1}$  from a three host plus one outflow component multi-Gaussian fit. As geometrical size we use  $R_{\text{out}} = 0.23''$  from data that are averaging  $v = [-330, -190] \text{ km s}^{-1}$  in the blue wing only, because the red wing is contaminated by nitrogen sulfide. Fluetsch et al. (2019) present CO(2-1) data for NGC 4418 with a line profile that is very similar to our NS-subtracted CO(1-0) profile (see also Fig. 2 below), but adopt a simple two-Gaussian decomposition that assigns a larger fraction of the total line flux to an outflow that dominates at  $|v| \gtrsim 100 \text{ km s}^{-1}$ . We maintain our more conservative decomposition but note that none of the available data exclude that such a large fraction of low-velocity emission is due to slow outflow-like motion, rather than gravitationally bound motion.

**IC 860 (Fig. A.8).** CO emission shows a prominent compact core with an intrinsic diameter of about  $1.5''$ . We find no significant emission at  $|v| > 250 \text{ km s}^{-1}$ . The  $|v| \lesssim 200 \text{ km s}^{-1}$  line wings near the nucleus are broader than the velocity of larger-scale rotation and the positions of these channels fold back to near the nucleus rather than extending the PA  $\sim 20$ – $30^\circ$  rotation pattern. We cannot discriminate some modest outflow motion (unable to escape the galaxy) from bound motion in the near-nuclear potential. We assign a limit of  $0.58 \text{ Jy km s}^{-1}$  to a fast outflow.

**IRAS 13120–5453 (Fig. A.9).** This merger shows prominent CO emission (see also Sliwa et al. 2017). Our data indicate a clear rotating disk extending to at least  $4''$ , with a projected rotation velocity of  $\sim 150 \text{ km s}^{-1}$ . Channel maps near systemic velocity show a depression at the center of this disk. The central line profile is double-horned with peak velocities similar to the larger-scale rotation, and shows additional outflow wings in the blue and red. Wings suggesting outflow are also seen in HCN(4-3) (Privon et al. 2017) and CO(3-2) (Fluetsch et al.

2019, see also Sect. 4.1). Spectral decomposition using four host plus two outflow Gaussian components results in the outflow Gaussians contributing roughly half of the  $r < 1''$  CO flux (15% of the total flux), if integrating over the full profile. This type of profile decomposition is backed by the need for a strong optically thin central outflow component in order to explain the relation of CO(3-2) and (1-0) emission in this source (see discussion in Sect. 4.1 below, including Figs. 2 and 3). We adopt  $v_{\text{out}} = 301 \text{ km s}^{-1}$  from this fit to the  $r = 1''$  outflow spectrum and  $R_{\text{out}} = 0.63''$  from averaging the properties of the blue and red outflows at  $v = [-690, -270] \text{ km s}^{-1}$  and  $v = [270, 690] \text{ km s}^{-1}$ . The adopted  $\Delta v$ ,  $FW_{10\%}$  and  $v_{\text{out}}$  are the flux-weighted means of the respective values for the two broad components.

**IRAS F14378–3651 (Fig. A.10).** The data show CO emission extended over a radius of  $\sim 2''$ , with a velocity gradient along PA  $\sim 20^\circ$ . The nuclear line profile is well fitted with two Gaussians. While there may be a slight flux excess over this fit at  $|v| \sim 250 \text{ km s}^{-1}$ , we do not consider this clear evidence for an outflow and assign a flux limit of  $0.43 \text{ Jy km s}^{-1}$  to a faster outflow.

**IRAS F17207–0014 (Fig. A.11).** The inner arcsecond of this merger shows a disk-like velocity gradient along PA  $\sim 50^\circ$  but larger scales do not follow this trend. The spatial and velocity structure of CO in this merger is complex, including a tidal arm that is detected in CO out to at least  $15''$ . The line profile of the galaxy is extremely broad and reaches out to  $\sim \pm 500 \text{ km s}^{-1}$ . Beyond that, an outflow is detected on the blue side only, out to  $v \gtrsim -1000 \text{ km s}^{-1}$ . An outflow covering a similar velocity range has been reported for CO(2-1) (García-Burillo et al. 2015, see also Sect. 4.1). Our lower-spatial-resolution data are consistent with their PA  $\sim 160^\circ$  for the blue outflow. We adopt  $v_{\text{out}} = 1116 \text{ km s}^{-1}$  from the fit to the nuclear outflow spectrum and an uncertain  $R_{\text{out}} = 0.45''$  from the properties of the blue outflow at  $v = [-1000, -500] \text{ km s}^{-1}$ .

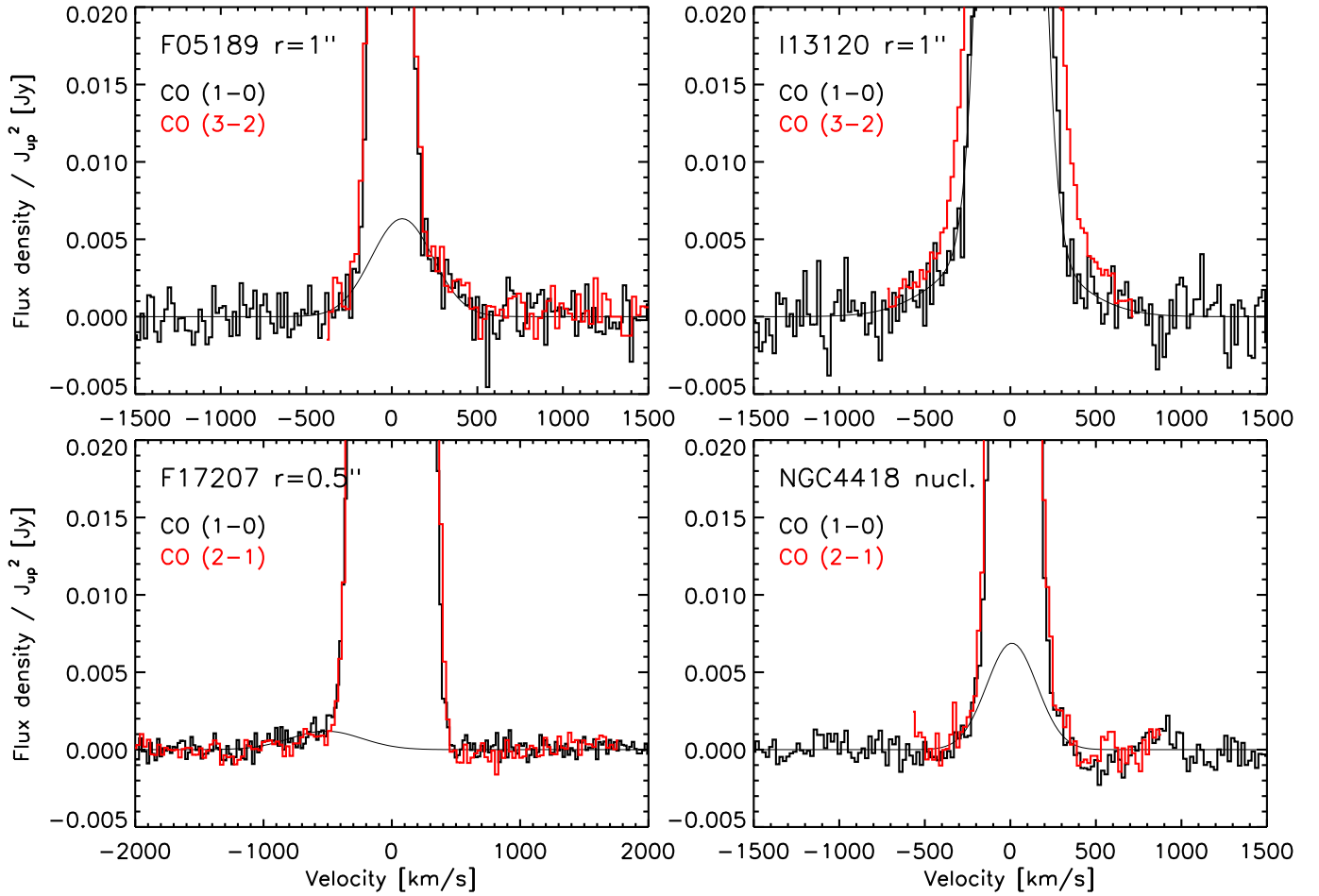
**IRAS F20551–4250 (Fig. A.12).** The velocity field of this merger is complex with a general NE–SW velocity gradient in the inner arcsecond. A clear red outflow is detected out to  $v \sim 500 \text{ km s}^{-1}$ , spatially offset towards the opposite direction as the red "disk" emission. A weaker blue counter-outflow is indicated at  $>4\sigma$  in the contours for the  $[-510, -210] \text{ km s}^{-1}$  range as shown in Fig. A.12, but only  $\sim 3\sigma$  for ranges ending at  $-270 \text{ km s}^{-1}$ . We base quantitative analysis on the red outflow only, but remain aware of the tentative blue outflow when comparing to other outflow indicators. We adopt  $v_{\text{out}} = 490 \text{ km s}^{-1}$  from the spectral fit at the position of the red outflow, and  $R_{\text{out}} = 0.27''$ .

**NGC 7479 (Fig. A.13).** There is strong CO(1-0) emission near the center and along the bar of this galaxy. The spectrum inside  $r = 3''$  shows no evidence for fast outflow at  $|v| > 300 \text{ km s}^{-1}$ . The centroid positions of different channel maps as well as the PV-diagram indicate a pronounced bipolar CO structure at  $|v| \sim 200 \text{ km s}^{-1}$ , at PA  $\sim 30^\circ$ . We do not ascribe this structure to outflow, but follow previous works (e.g., Sempere et al. 1995; Laine et al. 1999; Baker 2000), which on the basis of shallower data ascribed CO kinematics in the inner region of NGC 7479 to gravitationally bound gas flows within the barred potential of this galaxy. We assign a limit of  $0.73 \text{ Jy km s}^{-1}$  to a fast outflow.

### 3.2. Outflow data from the literature

We have searched the literature for local galaxies with interferometric low-J CO outflow detections or limits, starting with





**Fig. 2.** Excitation of outflowing molecular gas, comparing our CO(1-0) results (black, the fitted outflow component is also indicated) with higher-J results from [García-Burillo et al. \(2015\)](#) and [Fluetsch et al. \(2019\)](#). Spectra are divided by  $J_{\text{up}}^2$  so that they will overlap for excitation  $R_{ij} = 1$ , i.e., same brightness temperature.

the sources from Table 1 of [Lutz et al. \(2016\)](#) that we did not observe, but also including galaxies outside that parameter range. As noted in the introduction, our selection is not complete towards some of the smallest-scale and least energetic outflows that were recently reported from very-high-resolution interferometric studies. These would escape detection with “outflow wing” methods similar to ours, and be of lower impact beyond the local (nuclear) region of the host. Since different references follow a range of approaches for measuring outflow properties; we use the published results to reconstruct “observed”  $S_{\text{CO}}$  (wings),  $v_{\text{out}}$ ,  $R_{\text{out}}$  that are homogenized to our adopted conventions and consider differences in adopted distance and CO conversion factor where applicable. Where full parameters of multiple Gaussian fits are reported in the literature, our approach as described in Sect. 3 can be directly followed for deriving  $S_{\text{CO}}$  (wings) and  $v_{\text{out}}$ . For references that only quote outflow fluxes for a velocity range outside the host-dominated line core (e.g., [Cicone et al. 2014](#)), we adopt those fluxes and the average velocity of the outflow bin. If only the total  $S_{\text{CO}}$  (gauss) for a broad Gaussian component is available, we assume  $S_{\text{CO}}$  (wings) to be one-third of  $S_{\text{CO}}$  (gauss), where the factor of one-third is adopted from the typical value measured for our own data. Again, we do not attempt to correct for outflow opening angle and inclination. This provides a homogeneous sample but does not make full use of all the information at hand for some of the best studies cases (e.g., [Barcos-Muñoz et al. 2018](#); [Aalto et al. 2016](#)). In a few cases we use published plotted spectra to estimate  $v_{\text{out}}$ .

For 4C 12.50 we adopt the [Dasyra et al. \(2014\)](#) limit for CO outflow in emission but we add their outflow velocity from CO(3-2) absorption and their size estimate. For the single dish outflow detection of IRAS F17020+4544 and the low-resolution interferometry of IRAS 05083+7936 we adopt a fiducial 1 kpc for  $R_{\text{out}}$ . Our adopted homogenized literature data are summarized in Table 4.

From this point on, we treat our observational results and the adopted literature values jointly and consistently, as described in Sect. 4, and refer to this as the “combined” sample. For literature sources observed in CO transitions other than CO(1-0), we adopt the brightness temperature ratios in Table 4 of [Bothwell et al. \(2013\)](#) for conversion to CO(1-0).

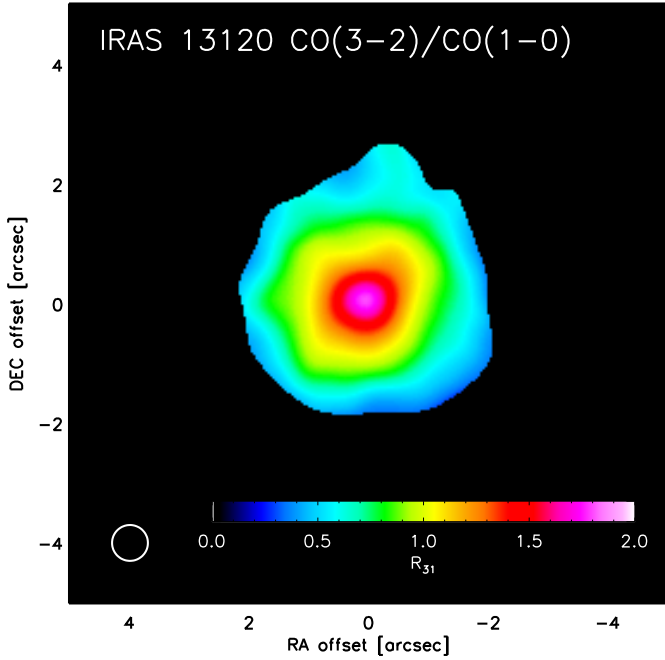
## 4. Outflow structure and derived properties

Deriving outflow masses  $M_{\text{out}}$ , mass outflow rates  $\dot{M}_{\text{out}}$ , and energy and momentum flows from the direct observed quantities (Sect. 3, Tables 3 and 4) is a multi-step procedure. We describe here our assumptions, and to what extent our data constrain the uncertainties in the conversions involved.

### 4.1. Outflow masses and CO conversion factor

Outflowing molecular gas may not necessarily be in self-gravitating molecular clouds for which the standard “Galactic” conversion factor from CO luminosity to molecular gas mass is





**Fig. 3.** Brightness temperature ratio  $R_{31}$  of CO(3-2) and CO(1-0) emission for the  $\pm 200 \text{ km s}^{-1}$  line core of IRAS 13120–5453. Data have been matched to the same beam. The nuclear region exceeds the maximum of 1 for optically thick thermal emission.

originally derived (see, e.g., Bolatto et al. 2013a). A range of possible conversion factors have been discussed for molecular outflows, ranging from a Galactic  $\alpha_{\text{CO}(1-0)} = 4.36 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  to  $\alpha_{\text{CO}(1-0)} = 0.80 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  as often used for ULIRGs (adopted partly due to the first outflow detections being in ULIRGs), and finally still lower values for  $\alpha_{\text{CO}(1-0)}$  that may arise in gas with optically thin CO emission. In that case,  $\alpha_{\text{CO}(1-0)} \approx 0.34 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  for an excitation temperature  $T_{\text{ex}} = 30 \text{ K}$  (Bolatto et al. 2013a), and the ratio of a rotational transition  $i$  and a lower transition  $j$  (in brightness temperature units) may exceed the value  $R_{ij} \lesssim 1$  for thermalized optically thick gas. Ratios may reach higher values up to  $R_{21} \approx 2.8$  and  $R_{32} \approx 1.3$ , for  $T_{\text{ex}} = 30 \text{ K}$  (e.g., Leroy et al. 2015a).

Indeed, high CO line ratios and optically thin CO have been suggested for the outflow of IC 5063 (Dasyra et al. 2016). For four sources in our sample we can compare our CO(1-0) outflow results to published spectra for higher-J CO transitions, which we have re-extracted for consistent apertures (Fig. 2). Values consistent with  $R_{i1} \lesssim 1$  are found for three targets:  $R_{31} = 1.35 \pm 0.23$  for IRAS F05189–2524 (where we combine with CO(3-2) from Fluetsch et al. 2019),  $R_{21} = 0.66 \pm 0.15$  for IRAS F17207–0014, (CO(2-1) from García-Burillo et al. 2015), and  $R_{31} = 1.34 \pm 0.23$  for NGC 4418 (CO(2-1) from Fluetsch et al. 2019). We derived these brightness temperature ratios by fitting multiple Gaussians to the higher-J transition where mean velocity and width of the outflow component(s) are fixed to the CO(1-0) derived values, and only their normalization is varied. The error estimates combine statistical errors and an assumed 10% calibration uncertainty for each line. These results do not point toward the low conversion factor of the optically thin case. This is consistent with similar ratios reported for other outflows (Weiß et al. 2005; Ciccone et al. 2012, 2018; Feruglio et al. 2015; Zschaechner et al. 2018), with  $\alpha_{\text{CO}(1-0)} = 2.1 \pm 1.2 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  suggested by Ciccone et al. (2018) on the basis of CO and [CI] in the N6240 outflow, and

with an attempt to constrain the outflow conversion factor for M82, leading to a  $\alpha_{\text{CO}(1-0)}$  that is two to four times below the Galactic value (Leroy et al. 2015a).

The clear exception in our sample is IRAS 13120–5453. The CO(3-2) spectrum (see also Fluetsch et al. 2019), divided by nine to get consistent brightness temperature scale with CO(1-0) clearly has wings above the CO(1-0) spectrum (Fig. 2). We derive  $R_{31} = 2.10 \pm 0.29$ , above the maximum value of 1 that is possible for thermalized optically thick CO emission, and consistent with a large contribution of optically thin gas that may have a moderate  $T_{\text{ex}} \sim 30 \text{ K}$ . This high excitation is not only observed in the line wings. Figure 3 shows the beam-matched ratio map of CO(3-2) and CO(1-0) flux (in brightness temperature units  $R_{ij}$ ), integrated over the line core within  $\pm 200 \text{ km s}^{-1}$ . Values corresponding to  $R_{31} \sim 0.6$  are observed in the disk, but  $R_{31} > 1$  is found in the nuclear region. Combined with the general CO morphology of the source (see moment-zero map and PV diagrams in Fig. A.9), the source is best described as a combination of a rotating disk with a radius of at least  $4''$  and normal “molecular cloud” CO properties, and a largely optically thin CO outflow that is dominating the nuclear region over the full line width.

Three of the four targets studied here and the vast majority of literature results are consistent with optically thick CO emission in the outflows. However, the clear detection of a high ratio (optically thin CO) for IRAS 13120–5453 is a strong warning that this is not true for all outflows, and that associated uncertainties on conversion factor and on outflow masses cannot be fully removed with the present data. Below, we follow much of the local universe molecular outflow literature and adopt  $\alpha_{\text{CO}(1-0)} = 0.8 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  but note that our CO data alone do not determine a preference for this “ULIRG” value compared to a Galactic one. For a Galactic  $\alpha_{\text{CO}(1-0)} = 4.36 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$ , recently found to be consistent also with ULIRG gas-to-dust comparisons (e.g., Tacconi et al. 2018), derived outflow masses and rates would increase more than five-fold. It is also important to bear in mind that outflows may be complex and multi-phase. For detection and current constraints from high-density tracers in the Mrk 231 and NGC 253 outflows, see Aalto et al. (2012a, 2015), and Walter et al. (2017).

In the absence of a reliable source-specific conversion factor, we also apply  $\alpha_{\text{CO}(1-0)} = 0.8 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  to IRAS 13120–5453, noting that the associated overestimate of outflow mass for the gas that is optically thin in CO and the underestimate due to including only line wings may partly compensate. Concerning the full sample, the good agreement that we report below with OH-based outflow masses (Sect. 5.1) is sensitive to our adopted CO conversion factor.

#### 4.2. Outflow rates

Given an observed outflow mass  $M_{\text{out}}$ , radius  $R_{\text{out}}$  and velocity  $v_{\text{out}}$ , mass-outflow rates are commonly derived by assuming a spherical or (multi-)cone geometry with constant velocity within that volume, and with the simplifying assumption of no temporal or spatial acceleration or deceleration of the flow. This is a nontrivial assumption given that a multitude of mechanisms such as ballistic deceleration in the potential of the galaxy, acceleration by radiation pressure, or interaction with an entraining nonmolecular wind can all affect velocities, but it is appropriate for interpretation of the current body of observations which is typically not highly resolved. The outflow rate is then derived via

**Table 4.** Adopted CO(1-0) outflow parameters for literature sources.

Source	$J_{\text{up}}$	$S_{\text{CO}}$ (wings) (Jy km s <sup>-1</sup> )	$v_{\text{out}}$ (km s <sup>-1</sup> )	$R_{\text{out}}$ ( $''$ )	Reference
NGC 253	1	707.00	50	13.35	Bolatto et al. (2013b)
IZw 1	1	<0.30			Cicone et al. (2014)
PG 0157+001	3	0.86	402	0.23	Fluetsch et al. (2019)
NGC 1068	3	98.00	100	1.50	García-Burillo et al. (2014)
NGC 1266	1	26.70	322	2.75	Alatalo et al. (2011)
IRAS F03158+4227	1	1.11	876	2.60	Gowardhan et al. (2018)
NGC 1377	2	29.77	138	1.96	Aalto et al. (2012b)
NGC 1433	3	16.10	100	2.00	Combes et al. (2013)
NGC 1614	1	3.55	360	1.73	García-Burillo et al. (2015)
IRAS F05081+7936	1	0.90	800	0.96	Leroy et al. (2015b)
NGC 2146	1	1450.0	200	25.00	Tsai et al. (2009)
IRAS F08572+3915	1	2.91	800	0.73	Cicone et al. (2014)
M 82	1	2413.33	100	53.00	Walter et al. (2002)
IRAS F10035+4852	1	<1.80			Leroy et al. (2015b)
NGC 3256 N	1	7.65	700	2.40	Sakamoto et al. (2014)
NGC 3256 S	1	5.65	500	4.00	Sakamoto et al. (2014)
IRAS F10565+2448	1	3.22	450	1.30	Cicone et al. (2014)
IRAS F11119+3257	1	0.40	1000	2.22	Veilleux et al. (2017)
NGC 3628	1	82.00	90	11.00	Tsai et al. (2012)
ESO 320-G030	2	5.21	340	3.71	Pereira-Santaella et al. (2016)
IRAS F12112+0305 SW	2	0.57	324	0.55	Pereira-Santaella et al. (2018)
IRAS F12112+0305 NE	2	10.70	465	1.15	Pereira-Santaella et al. (2018)
Mrk 231	1	4.16	700	0.72	Cicone et al. (2012)
Mrk 273	1	3.11	650	0.73	Cicone et al. (2014)
4C 12.50	1	<0.18	640	0.07	Dasyra & Combes (2012), Dasyra et al. (2014)
SDSSJ135656.10+102609.0	3	0.99	500	0.14	Sun et al. (2014)
IRAS F14348-1447 SW	2	7.60	419	1.05	Pereira-Santaella et al. (2018)
IRAS F14348-1447 NE	2	1.48	373	0.60	Pereira-Santaella et al. (2018)
Zw 049.057	2	5.00	300	0.30	Falstad et al. (2018)
Arp 220	1	0.88	380	0.20	Barcos-Muñoz et al. (2018)
Mrk 876	1	3.75	1050	1.54	Cicone et al. (2014)
NGC 6240	1	17.80	450	1.33	Feruglio et al. (2013)
IRAS F17020+4544	1	1.37	1670	0.86	Longinotti et al. (2018)
IRAS F17132+5313	1	<1.50			Leroy et al. (2015b)
PDS 456	3	1.42	900	1.00	Bischetti et al. (2019)
NGC 6764	1	1.40	53	7.00	Leon et al. (2007)
IRAS F20100-4156	1	1.12	929	1.00	Gowardhan et al. (2018)
IC 5063	2	12.30	500	2.00	Dasyra et al. (2016)
IRAS F22491-1808	2	2.00	325	0.35	Pereira-Santaella et al. (2018)
IRAS F23060+0505	1	2.26	700	1.38	Cicone et al. (2014)
IRAS 23365+3604	1	0.86	450	1.00	Cicone et al. (2014)

**Notes.** These parameters may differ from those published in the original references. They are reconstructed from the published information in order to be consistent with our adopted separation between line core and line wings, and our definitions for  $v_{\text{out}}$  and  $R_{\text{out}}$ . For consistency, they are not corrected for outflow inclination, even if this is constrained in some original references. See also Sect. 3.2 for details.

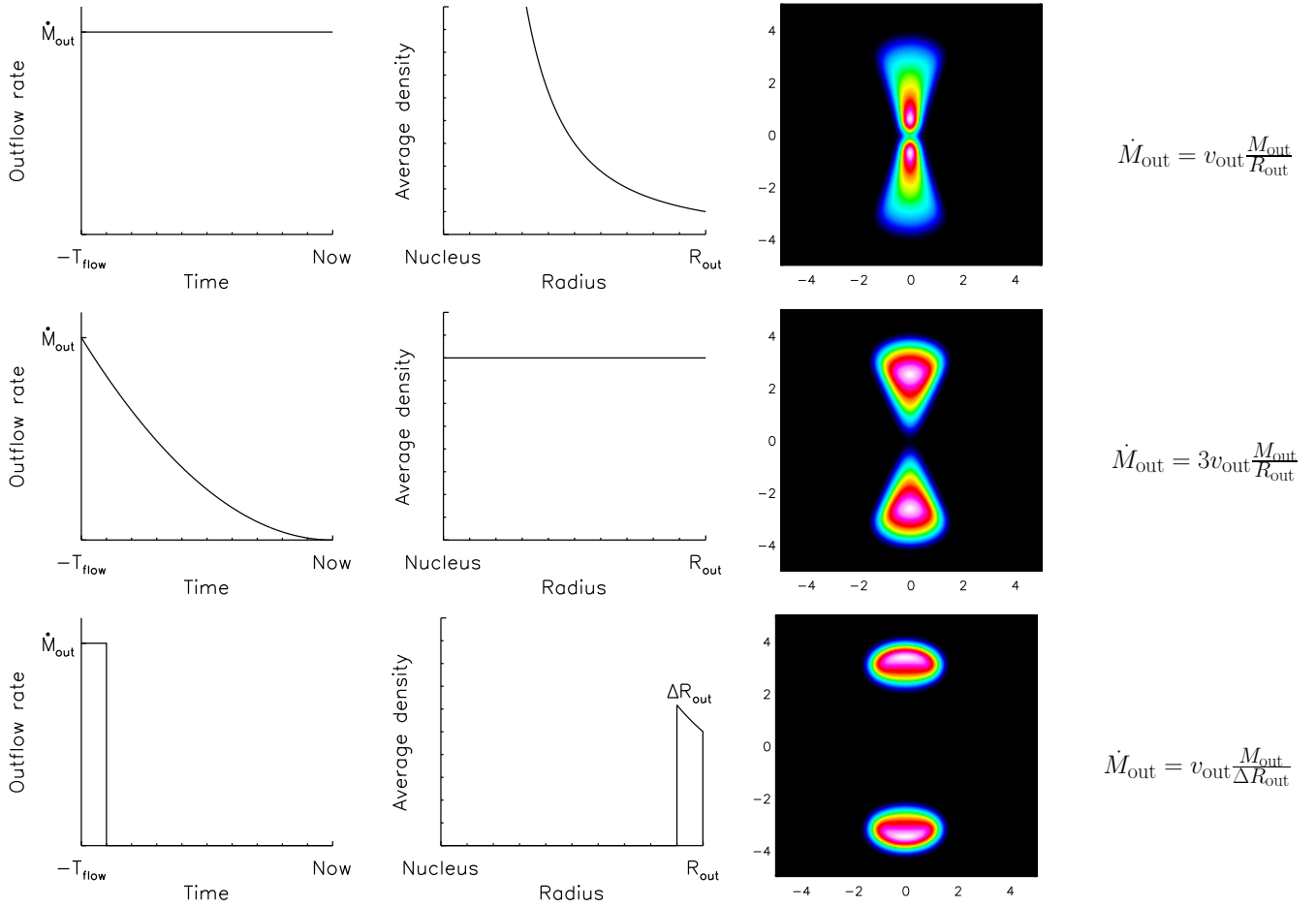
$$\dot{M}_{\text{out}} = C \frac{M_{\text{out}} v_{\text{out}}}{R_{\text{out}}}. \quad (3)$$

The value of  $C$  depends on the adopted outflow history (see also illustration in Fig. 4). Assuming that the outflow has started at a point in the past at  $-t_{\text{flow}} = -R_{\text{out}}/v_{\text{out}}$  and has continued with a constant mass outflow rate until now directly leads to  $C = 1$ . This is what we adopt, and agrees with the “time averaged thin shell” approach often used to derive outflow rates from absorption line data (e.g., Rupke et al. 2005; González-Alfonso et al. 2017). In this case, the average volume density of the outflowing gas decreases outwards  $\rho \propto R^{-2}$

but local cloud or filament densities may not follow this drop.

Several recent works have derived molecular outflow rates assuming  $C = 3$  (e.g., Cicone et al. 2014; Fiore et al. 2017). This is obtained when one assumes that the outflowing gas fills the spherical or (multi-)cone volume with constant average volume density. This corresponds to a mass outflow rate that starts with its maximum in the past at  $-t_{\text{flow}}$ , when the material was emitted that is now reaching  $R_{\text{out}}$ , and for which  $\dot{M}_{\text{out}}$  is quoted. After the start at  $-t_{\text{flow}}$  the mass outflow rate drops in time  $\propto (t/t_{\text{flow}})^2$ , now reaching zero.

Adopting a fixed CO conversion factor means CO light traces molecular gas mass. This is plausible if the outflow is seen as



**Fig. 4.** Schematic visualization of three commonly adopted outflow histories and their effects on radial density profile and conversion from outflow mass to outflow rate. *From left to right panels:* (i) outflow history, (ii) average radial density profile (local cloud conditions may differ from this average), (iii) a simulated moment-zero map for a modest  $20^\circ$  half opening angle flow seen with flow axis inclined by  $60^\circ$  with respect to the line of sight, and (iv) the equation for conversion from outflow mass to outflow rate. *From top to bottom:* (a) constant outflow history, (b) constant average volume density in cone, requiring a decaying outflow history, and (c) thin outflowing shell.

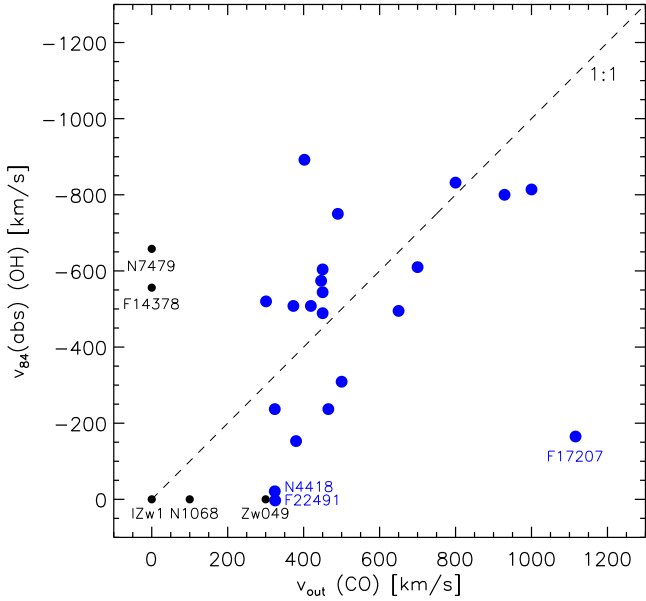
an ensemble of flowing clouds or filaments with conversion factor reflecting local cloud conditions rather than the larger-scale average density of the flow. This is an important assumption for both cases, but in particular for the first case ( $C = 1$ ) with average density decreasing outwards.

Both scenarios are strongly simplified, but for resolved outflows that are strongly inclined with respect to the line of sight and have moderate opening angle, the radial profiles can provide insights into the outflow history. Figure 4 schematically shows outflow history, radial profile of the average density, and projected surface brightness for these two scenarios and for a thin shell case. Most CO outflows in our sample and in the current literature are barely spatially resolved, preventing any related conclusions. An exception is III Zw 035 (Fig. A.1), the data for which suggest an inclined outflow of modest opening angle, lacking the brightening at large radius that is expected for the constant volume density or thin-shell cases (Fig. 4). Similar arguments can be made for the strongly inclined outflows of M 82 (Walter et al. 2002; Leroy et al. 2015a), NGC 253 (Bolatto et al. 2013b), and NGC 3256S (Sakamoto et al. 2014). For ESO 320–G030, substructure is observed, but again no brightening at large radii (Pereira-Santaella et al. 2016). We therefore proceed by computing outflow rates with  $C = 1$  in Eq. (3).

## 5. Widespread molecular outflows

### 5.1. Consistency of *Herschel* OH-based and interferometric CO outflow detections

Samples of dozens of galaxies each have now been studied using the two complementary techniques for detecting molecular outflows in local galaxies: P-Cygni profiles of far-infrared OH lines enabled by the *Herschel* mission and CO line wings enabled by wide-band receivers at mm interferometers. Comparing results from these two approaches is demanded in the first place by their basic difference. Hydroxyl absorption probes columns along the lines of sight towards the far-infrared-emitting region, inducing a need for model derivations of covering factor and radial location under the assumption of overall spherical symmetry. Carbon monoxide emission traces all emitting gas, but only to the extent that it can be kinematically and/or spatially separated from the host CO emission. Both methods invoke conversions from observables to total outflowing molecular gas mass that include explicit conversion factors – OH abundance and  $\alpha_{\text{CO}(1-0)}$ , respectively. We compare the two approaches in two ways. First, we use our combined CO sample and a large set of OH data from the literature (Veilleux et al. 2013; Spoon et al. 2013; Falstad et al. 2015; Stone et al. 2016) to check for consistency of detections and of the measured outflow velocities. Second, we compare the



**Fig. 5.** Comparison of outflow velocities measured from CO mm interferometry and from OH far-infrared P-Cygni profiles. For objects without outflow detection in one of the tracers, we use black symbols and arbitrarily set  $v$  to zero for the tracer without outflow detection.

outflow properties that [González-Alfonso et al. \(2017\)](#) derive for a dozen well-characterized OH outflow sources with CO results – interferometric data are available for their entire sample from our observations and from the literature.

Figure 5 compares CO- and OH-derived outflow velocities.  $v_{\text{out}}(\text{CO})$  is derived as described in Sect. 4. Our  $v_{\text{out}}(\text{CO})$ , which is not corrected for outflow inclination, is compared with the OH absorption in the line of sight. For OH, several different conventions are in use. We adopt  $v_{84}(\text{abs})$  as defined by [Veilleux et al. \(2013\)](#), since it is more robust at limited S/N than the alternative  $v_{\text{max}}$ , and better matched to our  $v_{\text{out}}$  convention for CO which samples the outflow wings but not the maximum detected velocity. In one case where the original OH reference quotes only  $v_{\text{max}}$  (IRAS F20100–4156, [Spoon et al. 2013](#)), we estimate  $v_{84}(\text{abs}) \sim 800 \text{ km s}^{-1}$  from their Fig. 10.

The agreement between the two approaches is good: Of 26 sources with data from both methods, 21 show both CO outflow and absorption dominated or P-Cygni OH spectra with available  $v_{84}(\text{abs})$ . The CO- and OH-based velocities are in broad agreement, but there is considerable  $\sim 0.36$  dex scatter around the 1:1 line. Correlations of outflow velocities with galaxy properties will be subject to related scatter, likely for either of the two methods in use.

Some insights can be gained also from the few galaxies for which CO and OH disagree in reported outflows. NGC 7479 shows a clear fast OH outflow with  $v_{84}(\text{abs}) \sim -650 \text{ km s}^{-1}$  ([Stone et al. 2016](#)) but no detectable CO outflow at such velocities (Fig. A.13). The line of sight sampling the OH outflow against the far-infrared continuum must have conditions that cannot be extrapolated to a larger scale and large-covering-factor outflow. Conversely, NGC 1068 exhibits no OH outflow ([Stone et al. 2016](#)) but molecular outflows have been suggested in CO studies. This discrepancy may be due to its moderate-velocity CO outflow occurring mostly in the disk plane (e.g., [García-Burillo et al. 2014](#)), and not affecting our line of sight. The faster outflow reported by [Gallimore et al. \(2016\)](#) might be closer to the line of sight, but restricted to a much smaller volume and covering only a small fraction of the far-infrared

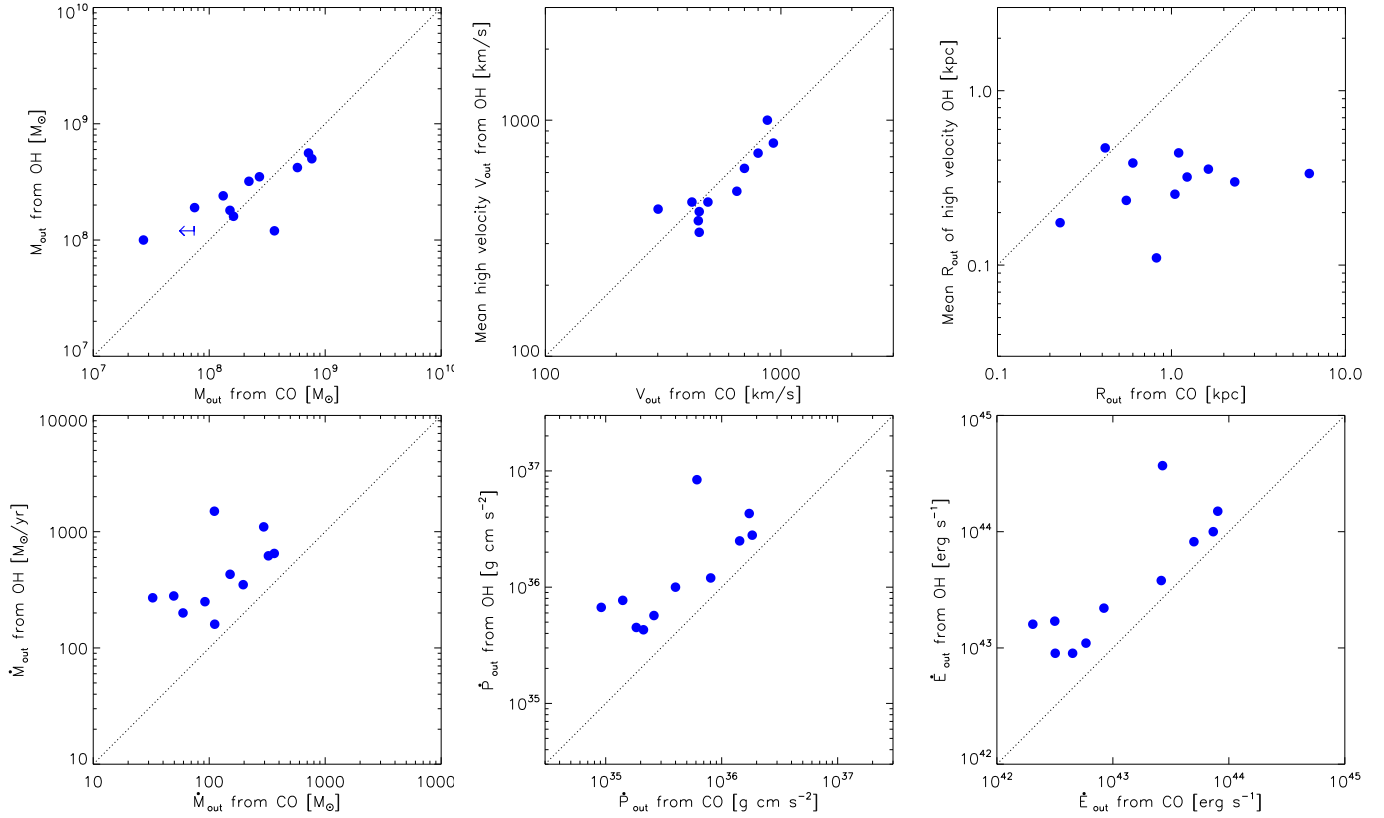
continuum. For NGC 4418, [González-Alfonso et al. \(2012\)](#) and [Veilleux et al. \(2013\)](#) report OH absorption but the absorption is slightly redshifted, suggesting molecular gas in slow,  $\sim 100 \text{ km s}^{-1}$  inflow. Inflow is also suggested for atomic gas by an inverse P-Cygni profile in [OI]  $63 \mu\text{m}$  ([González-Alfonso et al. 2012](#)). Any CO emission from this slow inflow component would be hard to separate from the line core in Fig. A.7. Finally, the modest OH outflow velocity for IRAS F17207–0014 clearly differs from the fast (but low flux density) outflow seen both in CO(1-0) (this paper) and CO(2-1) ([García-Burillo et al. 2015](#)), indicating that the OH column at these fast velocities must be low towards our line of sight.

As in the explanation of these mismatches in OH and CO outflow detection, we believe that for the comparison of OH and CO velocities of the outflows in Fig. 5 that are detected in both tracers, systematic factors and geometry are main contributors to the scatter. Similar arguments apply to the comparisons of outflow masses, velocities, and rates below in Fig. 6. Concerning CO outflow velocities, the nominal errors of typically  $\lesssim 10\%$  for the fitted width of the broad CO outflow component do not consider the effects of possibly non-Gaussian true profiles, nor the impact of varying flux and width of host CO emission on the splitting off of the outflow wings that are adopted in our analysis (Fig. 1). Concerning the OH  $v_{84}(\text{abs})$ , [Veilleux et al. \(2013\)](#) estimate a typical uncertainty of  $50 \text{ km s}^{-1}$  that is dominated by continuum placement uncertainties. Both estimates do not consider the different responses of the CO emission and OH absorption methods to variations in alignment of complex outflow geometries with respect to the line of sight, among other sources of systematic error such as the properties of the background far-infrared source against which OH is observed.

For the vast majority ( $\sim 80\%$ ) of the objects with both OH and CO diagnostics in hand, both OH absorption and CO outflow are detected, with plausible, mainly geometrical explanations for the remaining mismatches. For a sample of obscured nuclei characterized by HCN mm emission from vibrationally excited levels, [Falstad et al. \(2019\)](#) discuss in more detail geometrical and evolutionary factors that might lead to absence of far-infrared OH outflow signatures. The different sampling of the outflow geometry by emission and by absorption, as well as the modest S/N of many detections certainly also contribute to the scatter around the 1:1 line in the velocity comparison of Fig. 5. This good agreement would be too optimistic if dominated by CO followup of OH outflows, with preferential reporting of detections in the literature. However, CO detections and limits for our high-far-infrared surface brightness sample, and the analysis of the [González-Alfonso et al. \(2017\)](#) sample in the following paragraph, argue that such effects are not dominant.

A deeper comparison of CO-based outflow masses, rates, and energetics is possible for the 12 ULIRGs for which [González-Alfonso et al. \(2017\)](#) were able to derive reliable spherically symmetric outflow models that are constrained by both ground-state and excited level transitions of OH, observed with *Herschel*-PACS. All twelve sources are part of our combined CO sample. Figure 6 compares basic outflow mass, radius, and velocity as well as derived mass, momentum, and energy outflow rates. The OH models include up to three components per source; we follow [González-Alfonso et al. \(2017\)](#) in comparing CO data only to fast  $v \geq 200 \text{ km s}^{-1}$  OH components, since the CO equivalent of the slower OH components will be blended with the host line core. Masses are the sum of values for individual high-velocity OH components ( $v \geq 200 \text{ km s}^{-1}$ ), while radii and velocities are averages. Concerning outflow history, the assumptions of [González-Alfonso et al. \(2017\)](#) agree





**Fig. 6.** Comparison of CO-based outflow quantities with results from OH-based outflow models for 12 ULIRGs as derived by González-Alfonso et al. (2017). Dotted lines indicate a 1:1 relation. See also discussion in Sect. 5.1 on how this is influenced by differences in the regions sampled by OH absorption and CO emission.

with our use of  $C = 1$  in Eq. (3). For IRAS F14348–1447, Pereira-Santaella et al. (2018) report two distinct CO outflows from the two nuclei while the *Herschel* beam includes both. We use the dominant southwest CO outflow for comparison to OH.

Of the 12 ULIRGs, only IRAS F14378–3651 lacks a CO outflow detection and our assigned upper limit for CO-based outflow mass is within a factor of two of the OH result. In fact, Fig. A.10 indicates some excess at  $v \sim 250 \text{ km s}^{-1}$  above the fit with two Gaussians, but we have conservatively opted in Sect. 3 to not assign this excess at modest velocity to an outflow. Also, any weak CO flux excess at the velocities  $-500$  to  $-1000 \text{ km s}^{-1}$  which are reached by the OH outflow is not significant. For the remaining 11 ULIRGs, outflow masses agree very well (top left of Fig. 6), with mean  $\log(M_{\text{out,OH}}/M_{\text{out,CO}}) = 0.06 \text{ dex}$  and dispersion  $0.29 \text{ dex}$ . This agreement is noteworthy given the independent derivation of these masses, including in particular the conversion factor  $\alpha_{\text{CO}} = 0.8 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  on the CO side, and an interstellar medium OH abundance  $X_{\text{OH}} = 2.5 \times 10^{-6}$  on the OH side. A consistent lower limit on  $X_{\text{OH}}$  was obtained by Stone et al. (2018). The good agreement can be regarded as support for the adopted  $\alpha_{\text{CO}(1-0)}$ , but see also the discussion below in the context of outflow sizes.

Outflow velocities are in excellent agreement for these bright and well-studied galaxies (top center), with mean  $\log(v_{\text{out,OH}}/v_{\text{out,CO}}) = -0.03 \text{ dex}$  and dispersion  $0.08 \text{ dex}$ .

Larger but still modest differences are found for the outflow radii. Outflow radii from the OH models are around  $0.3 \text{ kpc}$  for all ULIRGs, while the CO observations vary more widely and detect radii up to several kiloparsecs. A mean  $\log(R_{\text{out,OH}}/R_{\text{out,CO}}) = -0.54 \text{ dex}$  is found with a dispersion of  $0.39 \text{ dex}$ . Similar  $\approx 0.5 \text{ dex}$  differences are found for the outflow

rates for mass, momentum, and mechanical energy (bottom panels of Fig. 6). While the OH radii are based on spherically symmetric models and the CO radii are measured, it is improbable that all differences stem from modeling uncertainties, especially when considering the observed OH absorptions in the  $84 \mu\text{m}$  and/or  $65 \mu\text{m}$  transitions that are radiatively excited by intense far-infrared radiation fields close to an optically thick and necessarily small far-infrared source (González-Alfonso et al. 2017). Breakup into several individually optically thick sources could enlarge the sizes but this is not supported by the usually compact structure of high-resolution (sub)mm images (e.g., Barcos-Muñoz et al. 2017).

Assuming that much of the size difference between OH and CO outflow (top right of Fig. 6) is real, one is led to a scenario where a more compact and embedded molecular outflow region traced by the OH absorptions connects to a larger and older component that can still be traced by the collisionally excited CO emission. In contrast, the covering by clumps absorbing OH will decrease with outward motion. In that case, the good agreement of outflow masses (top left of Fig. 6) will be partly fortuitous or erroneous since the two molecules trace different though overlapping regions. A further factor in this comparison arises from the definition of outflow masses. Hydroxyl-based outflow masses from González-Alfonso et al. (2017) reflect all directions of a spherically symmetric outflow. Emission line components of such a flow that are moving close to the plane of sky blend with the bright core of the line and are excluded by our definition of  $S_{\text{CO}}$  (wings); see also Fig. 1. Toy-model calculations for a spherically symmetric case suggest  $S_{\text{CO}}$  (wings) would measure a mass two to three times smaller than the  $4\pi$  outflow, but this correction is obviously dependent on outflow velocity and the width of the line core, and

may be an upper limit if replacing the spherically symmetric flow with a wide-angle flow with a large line-of-sight component.

A non-unique scenario that further reduces the small tensions in Fig. 6 between OH- and CO-based results can be described as follows: (i) An outflow that is constant in rate and velocity over the flow time derived from CO data is sampled in its entirety by CO emission, and its inner half or somewhat less is sampled by the OH absorptions. The molecular outflow masses traced by CO and by OH should reflect the different fractions of the flow that are sampled – the CO-based molecular outflow mass should be larger than the OH-based one. (ii) The similar derived outflow masses (Fig. 6) must therefore be due to an additional correction. Depending on actual outflow geometry, this could be achieved by considering the fraction of an approximately spherically symmetric CO outflow that is missed by the emission line wings (see above). Alternatively, increased  $\alpha_{\text{CO}(1-0)}$  or increased OH abundance (decreased molecular hydrogen mass traced by OH) could help in achieving an appropriate ratio of masses traced by CO and OH. (iii) Either of the corrections described in (ii) would also bring the outflow rates in the bottom panels of Fig. 6 closer to agreement, consistent with the assumed constant outflow rate.

In summary, the comparisons provide a very successful cross-validation of the OH P-Cygni and CO interferometric methods of characterizing molecular outflows. There is no indication for a need to modify the methods and conversions adopted for CO outflows in Sects. 3 and 4, nor for modifying the OH modeling methods of González-Alfonso et al. (2017). At least for ULIRG-like systems, which constitute the bulk of current outflow samples and of the comparisons above, the comparison also supports the large molecular outflow covering factor reported in OH outflow studies. This finding is originally based on the high incidence of P-Cygni absorptions in ULIRGs (e.g., Sturm et al. 2011; Veilleux et al. 2013), and is supported by the model results of González-Alfonso et al. (2017). Current knowledge of outflow geometries and histories for large samples is however insufficient to clarify whether remaining modest differences between OH- and CO-based results can find a purely geometrical explanation. Use of the time-efficient OH spectroscopy for high- $z$  outflow searches is encouraged, be it with ALMA (Spilker et al. 2018; Herrera-Camus et al. 2020a) or with planned future far-infrared/submm space facilities (e.g., SPICA, Roelfsema et al. 2018).

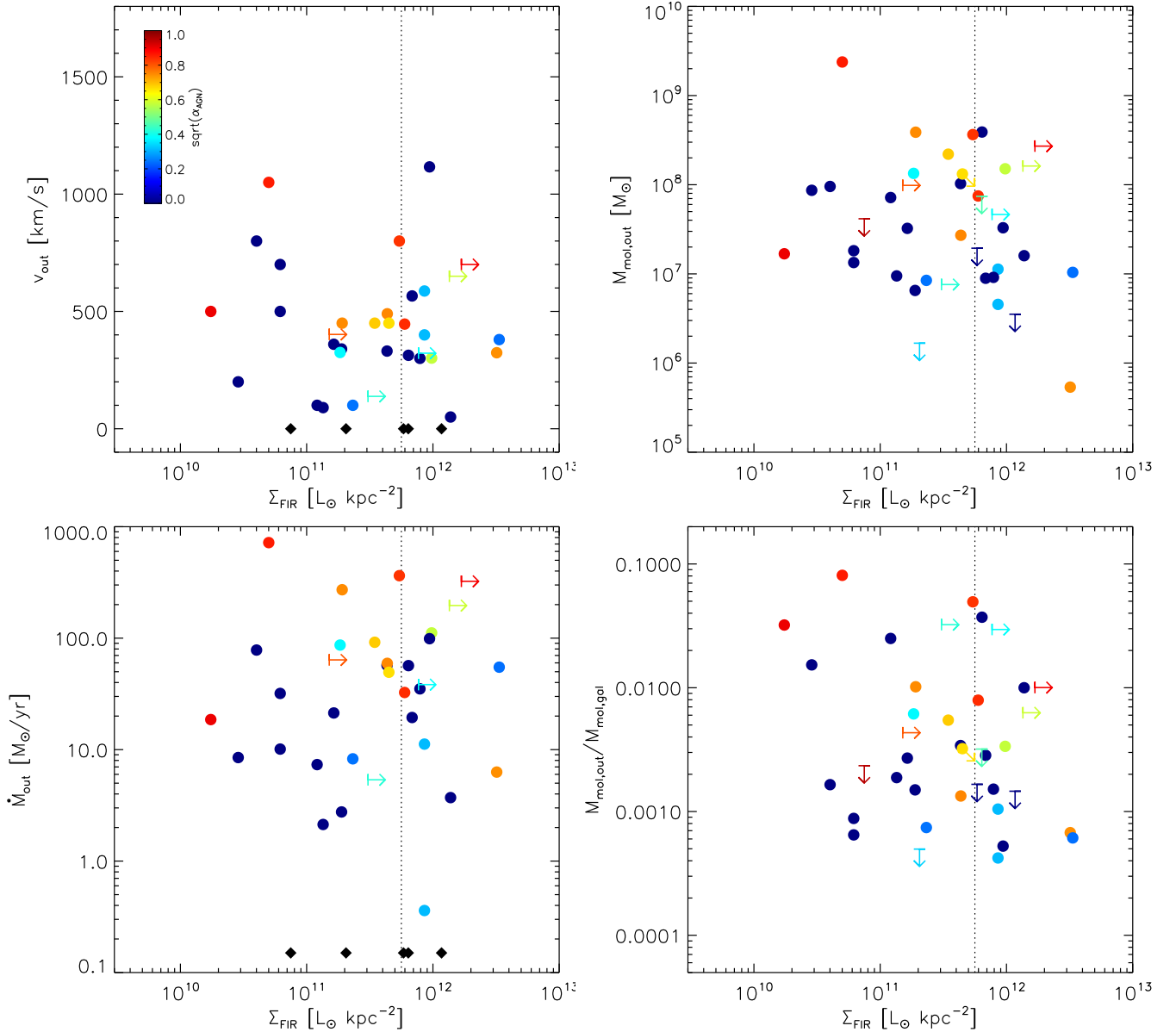
## 5.2. Relation of outflow properties to far-infrared surface brightness, AGN luminosity, and bolometric luminosity

Several recent papers have studied scaling relations of molecular outflow masses, mass outflow rates, momentum flows, and energy flows with star formation rate, AGN luminosity, and bolometric luminosity of the host galaxies (Cicone et al. 2014; Fiore et al. 2017; González-Alfonso et al. 2017; Fluetsch et al. 2019), and to what extent the flows may be momentum or energy driven. Here, we focus on the additional comparison with far-infrared surface brightness, and briefly revisit the relations to AGN and bolometric luminosity. Table C.1 summarizes these and related properties for the combined sample. As far as other quantities are concerned, our sample concurs with observational results from these references with modest modifications due to different assumptions (e.g.,  $C = 1$  or 3 in Eq. (3) and outflow line fluxes corresponding only to wings or to entire broad components). One should note that all these studies have overlapping samples, since they are enlarging their original samples from the literature (Sect. 3.2 for this paper), or are entirely literature based.

Figure 7 shows outflow velocity, outflow molecular mass, mass outflow rate, and the ratio of outflow and host molecular gas mass as a function of far-infrared surface brightness for 39 sources in our combined sample, which have measurements or limits for far-infrared surface brightness (Lutz et al. 2016, 2018). None of these quantities shows a significant trend with far-infrared surface brightness, but see the note below on selection effects at low surface brightness. If we first focus on the high-surface-brightness end with  $\log(\Sigma_{\text{FIR}}) \geq 11.75 L_{\odot} \text{ kpc}^{-2}$ , we indeed find the expected very high incidence of molecular outflows that motivated our observations. Of the 18 sources with  $\log(\Sigma_{\text{FIR}}) \geq 11.75 L_{\odot} \text{ kpc}^{-2}$  and/or close to optically thick far-infrared emission that are listed in Table 1 of Lutz et al. (2016), all but ESO 173–G015 have deep CO observations probing for outflows, from our own data or from the literature in the combined sample. Of these 17 objects, CO outflows have been detected in 13 sources while upper limits have been derived for IRAS F10173, IC 860, IRAS F14378, and NGC 7479. In several cases, these limits are conservative and a contribution of an outflow to the observed moderate velocity line wings remains a possibility. As expected, galaxies with high-far-infrared surface brightness are excellent hunting grounds for molecular outflows. This is reflected in the high detection rate: at least 13 out of 18 sources. On the other hand, the cases with sensitive upper limits demonstrate that far-infrared surface brightness alone is not sufficient to produce strong and fast outflows in all cases, despite strongly exceeding SFR thresholds above which outflows from galaxy disks are prevalent (Heckman 2002; Newman et al. 2012; Davies et al. 2019).

Given its selection above an already high surface-brightness threshold and the modest sample size, it is not surprising that no trends of outflow properties with surface brightness are detected within this high-surface-brightness sample. More noteworthy is the absence of trends in Fig. 7 over the full range of the combined sample, which extends to surface brightness that is more than an order of magnitude lower. A high incidence of molecular outflow is also found among the low-surface-brightness members of this sample. However, severe and poorly known selection effects may be at work here. Much of the molecular outflow literature is in papers studying single or very few sources, and selection criteria are diverse. Upper limits may often remain unpublished. We certainly cannot exclude the abundant existence of low-surface-brightness galaxies that have only weak or no molecular outflows. An obvious speculation is that the molecular outflow detection rate is smaller in this regime than for the high-surface-brightness sample. However, unbiased studies of larger samples are needed to establish such a difference, or to detect any trends in outflow properties over a large range of far-infrared surface brightness.

Figures 8 and 9 show the same outflow quantities for the combined sample, now in relation to AGN luminosity and to bolometric luminosity, as also listed in Table C.1. Consistent with previous results that are based on samples that are partly overlapping with our combined sample (Sturm et al. 2011; Veilleux et al. 2013; Spoon et al. 2013; Cicone et al. 2014; Fiore et al. 2017; González-Alfonso et al. 2017; Fluetsch et al. 2019), we find higher outflow velocities, outflow masses, and outflow rates for sources that host a luminous AGN. However, to be clearly noted in our larger sample (Fig. 8) is the presence of some quite important outflows also in sources with low values or upper limits for  $L_{\text{AGN}}$  (e.g., III Zw 035) that were not included in early work. There is no obvious link of outflow properties to  $L_{\text{AGN}}$  at low  $L_{\text{AGN}} \lesssim 10^{10} L_{\odot}$ ; the observed trends are driven by the high  $L_{\text{AGN}}$  sources only. Such increased scatter in the relation



**Fig. 7.** Relation of outflow properties to far-infrared surface brightness. The AGN contribution to the bolometric luminosity is color-coded on a square-root scale between zero (dark blue) and one (dark red). Black symbols in the *left panels*, plotted at an arbitrary low value, refer to sources lacking an outflow detection (upper limits in the *right panels*). High far-infrared surface brightness sources from Table 1 of [Lutz et al. \(2016\)](#) are at  $\log(\Sigma_{\text{FIR}}) \gtrsim 11.75$   $L_{\odot} \text{ kpc}^{-2}$  (vertical dotted lines). Carbon monoxide observations are close to complete here (17/18), unlike the incomplete literature results for lower surface brightness.

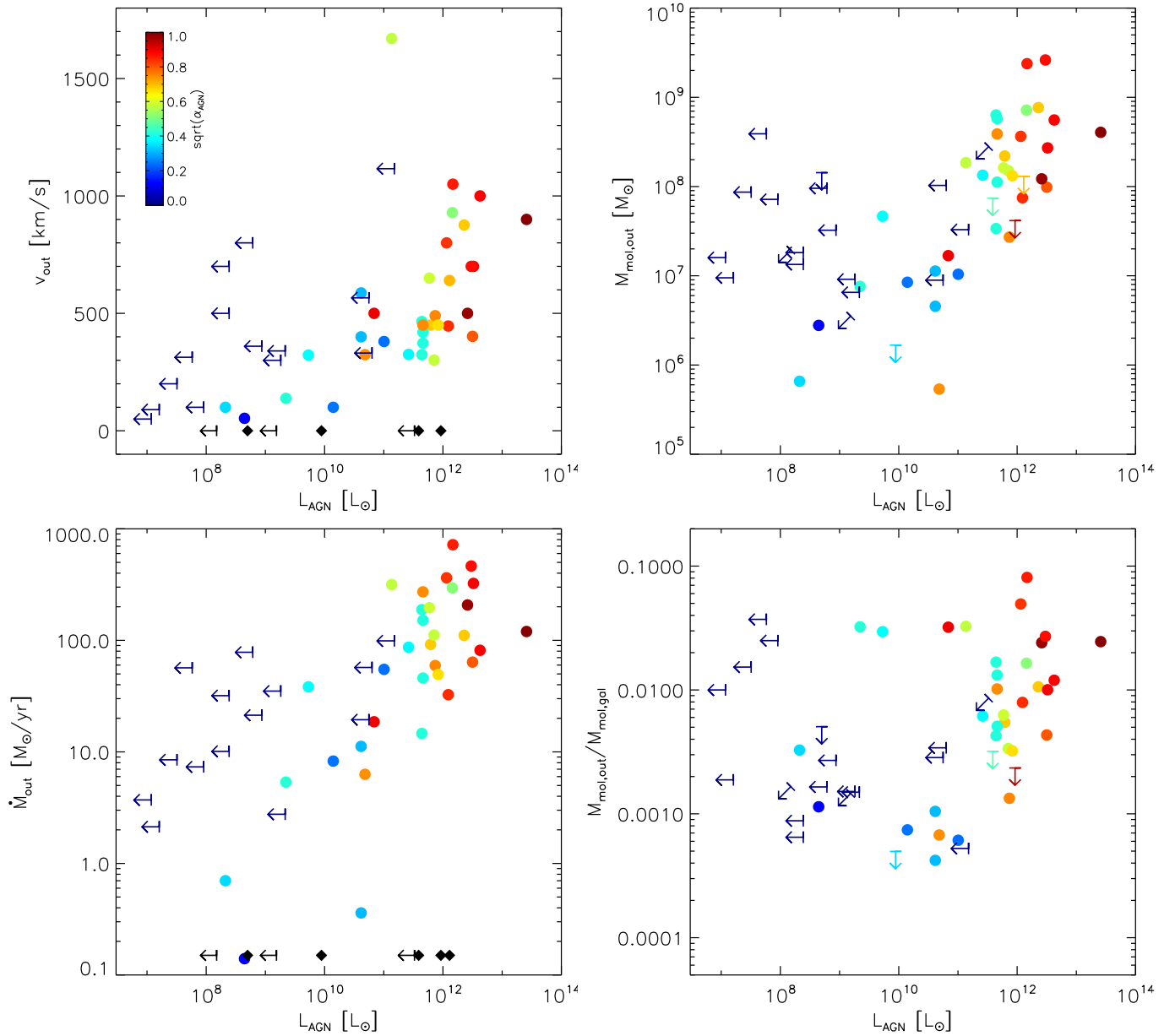
to AGN luminosity was also noted by [Fluetsch et al. \(2019\)](#). A further contribution to variation in outflow properties for a given  $L_{\text{AGN}}$  will arise from different geometries affecting the coupling of AGN wind to the larger-scale interstellar medium. In combination, these findings suggest an important though not exclusive role of driving that is linked to present AGN activity.

In fact, quite clear relations of molecular outflow mass and mass outflow rate with bolometric luminosity are observed (Fig. 9), which we quantify as:

$$\log(M_{\text{out}}/M_{\odot}) = (7.76 \pm 0.09) + (0.78 \pm 0.12) \times (\log(L_{\text{Bol}}/\text{erg s}^{-1}) - 45.5), \quad (4)$$

$$\log(\dot{M}_{\text{out}}/M_{\odot} \text{ yr}^{-1}) = (1.64 \pm 0.09) + (0.75 \pm 0.10) \times (\log(L_{\text{Bol}}/\text{erg s}^{-1}) - 45.5). \quad (5)$$

Here we used `linmix_err` ([Kelly 2007](#)) and adopted errors of 0.15 dex for the bolometric luminosities and 0.3 dex for the outflow quantities. The value of 0.15 dex is conservatively larger than uncertainties in measured  $L_{\text{IR}}$  or optical AGN continuum, in order to also consider uncertainties in the extrapolation from these to bolometric luminosity, and possible AGN variability. The fit for outflow mass is derived including both detections and upper limits, while the fit for mass-outflow rates includes only the outflow detections since  $R_{\text{out}}$  and  $v_{\text{out}}$  are needed to derive outflow rate or limit. We have verified that adopting fiducial  $R_{\text{out}} = 0.7$  kpc and  $v_{\text{out}} = 450$  km s $^{-1}$  to tentatively include the limits on outflow mass also for outflow rate leaves the slope of Eq. (5) effectively unchanged, and reduces the normalization by  $\sim 0.1$  dex. While these relations can be helpful in assessing the likely importance of a molecular outflow from a given object, or for estimating total outflow rates for ensembles, we caution again



**Fig. 8.** Relation of outflow properties to AGN luminosity. The AGN contribution to the bolometric luminosity is color-coded on a square-root scale between zero (dark blue) and one (dark red). Black symbols in the *left panels*, plotted at an arbitrary low value, refer to sources lacking an outflow detection (upper limits in the *right panels*).

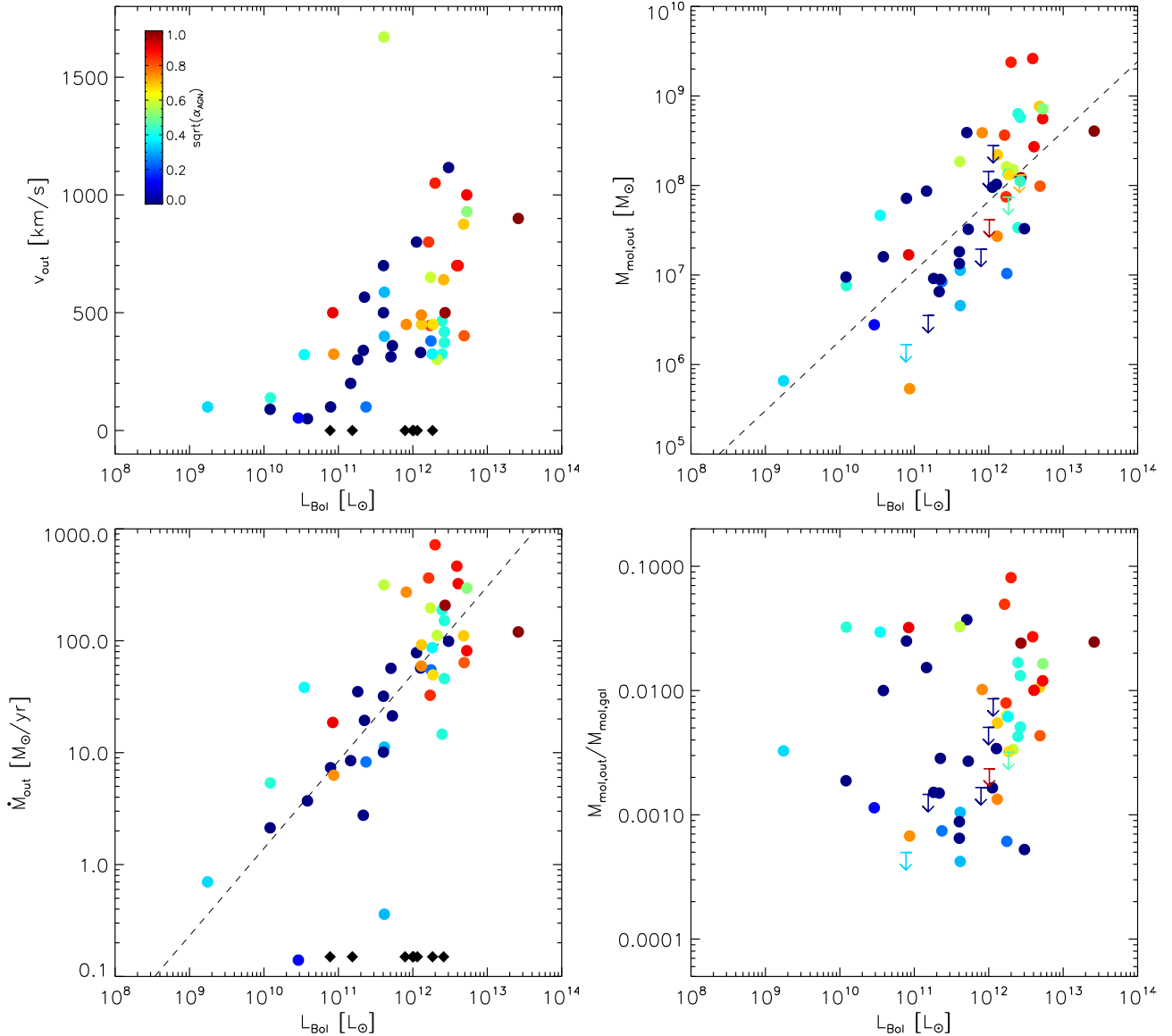
about possible selection effects in the combined sample that is significantly literature based. In addition, bolometric luminosity, AGN luminosity, and SFR are all correlated for this sample, causing a link of outflow properties to all three, including the SFR that we do not discuss in detail here.

There are two obvious ways to reconcile the preference of strong molecular outflows for AGN hosts with the detection of outflows also in galaxies that are lacking an AGN, and with the increased scatter at low  $L_{\text{AGN}}$  in Fig. 8. First, intense star formation is a significant contributor to driving molecular outflows, which becomes dominant in the absence of a powerful AGN. Second, AGN-driven outflows may not be recognized as such because they persist for some time even if rapid accretion-rate variations lower the AGN luminosity. A third possibility is driving by highly obscured AGN activity, but we note that for many cases the AGN luminosities in Table C.1 use infrared spectroscopy which is sensitive to AGN-heated dust even for

obscured cases. We believe that both star formation driving and intermittent AGN activity are at work here, and address the second option in some more detail in Sect. 5.3. First however we discuss constraints from kinetic power and momentum rate balance.

Figure 10 compares the outflow kinetic power  $P_{\text{kin}} = 0.5 \dot{M}_{\text{out}} v_{\text{out}}^2$  with AGN luminosity and bolometric luminosity. Most AGN hosts stay below a wind kinetic power of 5% of the AGN luminosity, which is often considered the limit for coupling AGN radiative power into outflows via “energy-conserving” winds (e.g., King & Pounds 2015). Three AGN hosts (NGC 1266, Mrk 876, IRAS F17020+4544) hover around or slightly above that line, still consistent with plausible uncertainties of wind parameters or with modest AGN variability. Notable is that limits on AGN luminosity place many of the sources with nondetected AGNs well above that line. Also, their outflow momentum rate would sometimes require “momentum boosts”





**Fig. 9.** Relation of outflow properties to bolometric luminosity. The AGN contribution to the bolometric luminosity is color-coded on a square-root scale between zero (dark blue) and one (dark red). Black diamonds in the *left panels*, plotted at an arbitrary low value, refer to sources lacking an outflow detection (upper limits in the *right panels*). Dashed lines reflect the fitting relations from Eqs. (4) and (5).

far above the maximum AGN radiative momentum that corresponds to the limit on  $L_{\text{AGN}}$ . This boost reaches up to more than a factor 100 (Fig. 11), beyond plausibility even for energy-conserving flows (Faucher-Giguère & Quataert 2012; Zubovas & King 2012). These two findings suggest that some of the observed molecular outflows cannot be driven by the currently observed level of AGN activity.

The right panels of Figs. 10 and 11 compare outflow power and momentum rate to bolometric luminosity and associated momentum rate, which for the nonAGN systems are both dominated by star formation. Comparing outflow kinetic power to the total mechanical luminosity delivered by supernovae  $P_{\text{kin,SN}} = 7 \times 10^{41}$  (SFR/ $M_{\odot} \text{ yr}^{-1}$ )  $\text{erg s}^{-1}$  (Veilleux et al. 2005), a relatively high conversion efficiency  $\geq 10\%$  would be needed in some cases. Similarly, the momentum rate of some of the stronger outflows from nonAGN hosts exceeds the radiative momentum rate. While this is in principle possible via multiple scattering in heavily obscured regions (e.g., Murray et al. 2005), we consider it more likely that we are facing a mix of driving by star formation and by past AGN activity.

Given the relatively clear relations of outflow kinetic power and momentum rate with bolometric luminosity, we quote fitting relations, again with a caveat about possible selection effects from the literature selection:

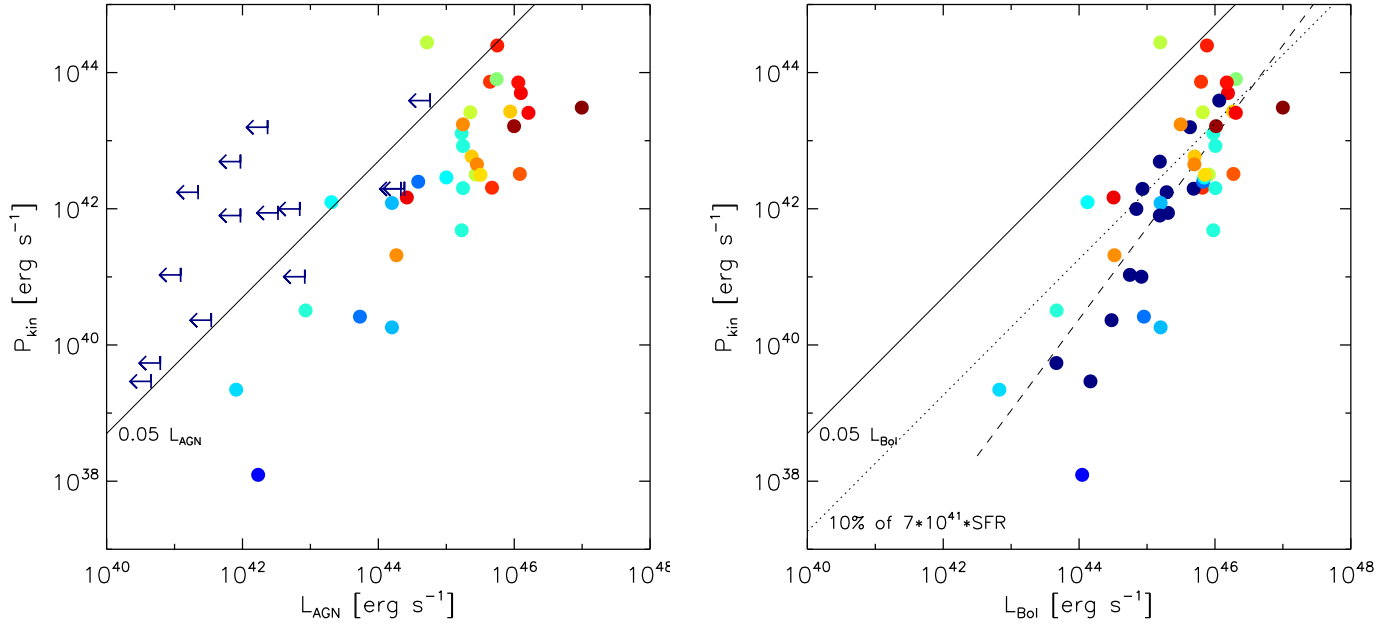
$$\log(P_{\text{kin}}/\text{erg s}^{-1}) = (42.39 \pm 0.14) + (1.34 \pm 0.17) \times (\log(L_{\text{Bol}}/\text{erg s}^{-1}) - 45.5), \quad (6)$$

$$\log(\dot{M}_{\text{out}} v/\text{g cm s}^{-2}) = (35.07 \pm 0.11) + (1.04 \pm 0.13) \times (\log(L_{\text{Bol}}/\text{erg s}^{-1}) - 45.5). \quad (7)$$

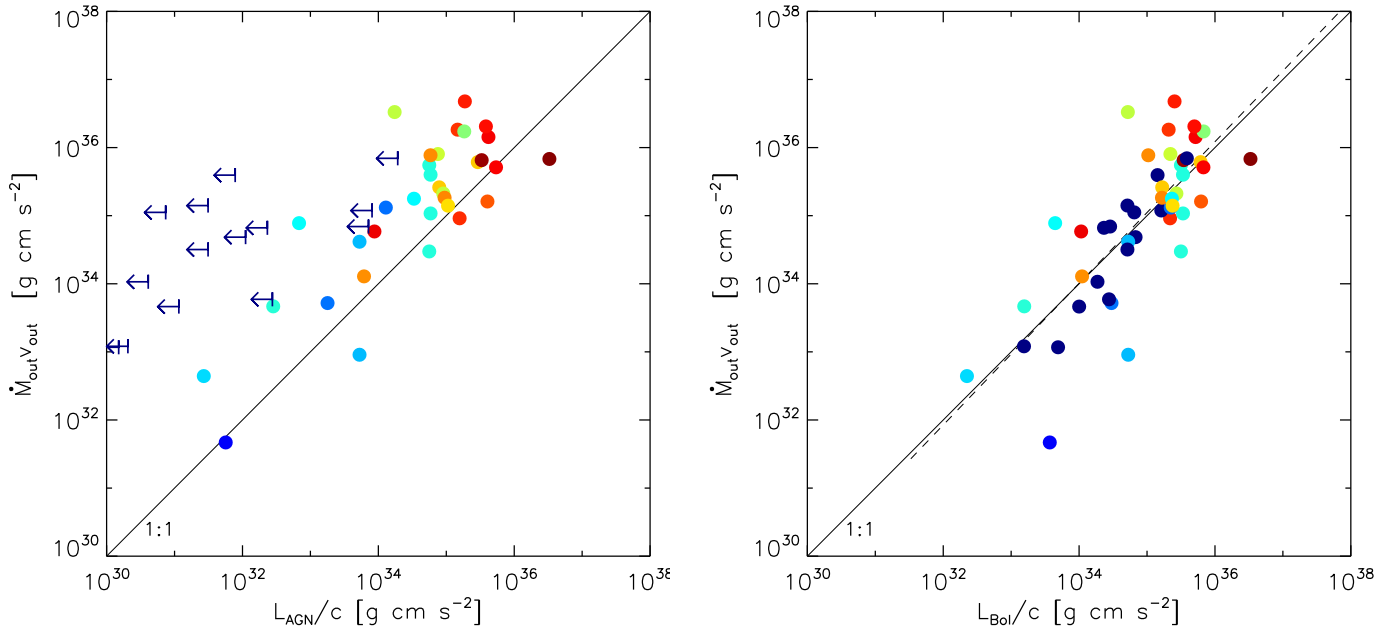
### 5.3. Evidence for a role of intermittent AGN driving

We now present arguments that the contribution of AGNs to the driving of outflows must have a short-timescale and intermittent component in order to match the observed outflow properties.

Figure 12 shows that typical flow times  $t_{\text{flow}} = R_{\text{out}}/v_{\text{out}}$  for the combined sample are short, just above a million years (median 1.5 Myr). For an outflow which is so massive that it needs to be driven by not only a single young cluster or



**Fig. 10.** Relation of outflow kinetic power to AGN luminosity (*left*) and bolometric luminosity (*right*). The AGN contribution to the bolometric luminosity is color-coded on a square-root scale between zero (dark blue) and one (dark red). The continuous line indicates coupling of 5% of the luminosity into outflow power, while the dotted line indicates coupling of 10% of the supernova kinetic power into outflow power (assuming the bolometric luminosity is SFR dominated). The dashed line in the right panel reflects the fit from Eq. (6).

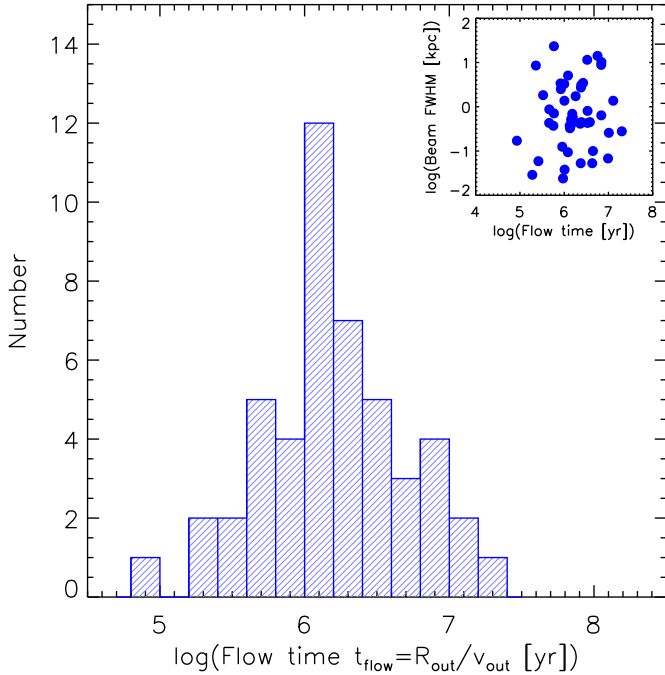


**Fig. 11.** Relation of outflow momentum rate to AGN luminosity momentum rate (*left*) and bolometric luminosity momentum rate (*right*). A 1:1 relation is over-plotted. The dashed line in the *right panel* reflects the fit from Eq. (7). The AGN contribution to the bolometric luminosity is color-coded on a square-root scale between zero (dark blue) and one (dark red).

star formation region but rather by a large part of a galaxy’s total (SFR+AGN) energetics and/or momentum, such a short timescale is difficult to achieve with star formation events. Given the star formation history and the smoothing effect of massive star lifetimes, star formation events would naturally lead to a  $\geq 10^7$  yr timescale, even in merger-induced starbursts with peaked star formation histories (e.g., Mihos & Hernquist 1996; Hopkins et al. 2008, 2013; Hayward et al. 2014). In a suite of 112 hydrodynamic major merger simulations from Cox et al. (2006) and Wuyts et al. (2010), the star formation rate stays

above half of its peak value for a median duration of  $\sim 70$  Myr, with a minimum of 5 Myr and only three cases with duration below 10 Myr. Short outflow events of  $\sim 1$  Myr in massive galaxies are therefore much easier to explain with the highly and rapidly variable accretion onto an AGN, even if considering that observed outflows may average to some extent over rapid repetitive AGN “flickering” (Zubovas & King 2016).

This argument assumes that the  $R_{\text{out}}$ ,  $v_{\text{out}}$  values obtained from CO interferometry and used to derive the flow time are representative for the full duration of the outflow event, have



**Fig. 12.** Histogram of flow times for the combined sample. The insert compares flow time with physical beam size.

constant velocity, and start at the center of the galaxy. This assumption would be incorrect if molecular gas from earlier phases of this outflow event were to remain undetected due to observational limitations, or due to being converted to atomic or ionized gas phases. What might happen to such molecular gas? If it is found at low surface density and CO surface brightness, observations with overly small beams might miss an important fraction of the outflow. While such effects surely exist, we infer from the lack of a trend in the inset of Fig. 12 that beam size effects are not dominating the observed flow-time distribution. Also, accumulated fast molecular gas from long outflow episodes has not been reported in large-aperture single-dish data that are sensitive to low-surface-brightness CO. In a situation where a large fraction of the outflowing gas will not finally escape the halo (e.g., [Fluetsch et al. 2019](#)), flow times as derived here are also limited by deceleration. Such effects will however work on  $\gtrsim 10^7$  yr timescales related to the dynamical timescale, longer than the short typical  $t_{\text{flow}}$  observed here.

Several studies have compared molecular outflows to outflows in the atomic phase and the ionized  $\sim 10^4$  K phase and concluded that, while the mass in those phases may sometimes approach the outflowing molecular mass, it does not exceed it by large factors ([Contursi et al. 2013](#); [Rupke & Veilleux 2013](#); [Leroy et al. 2015a](#); [Janssen et al. 2016](#); [Rupke et al. 2017](#); [Fiore et al. 2017](#); [Fluetsch et al. 2019](#)). Specifically for M82, [Leroy et al. \(2015a\)](#) suggest a molecular fraction in the outflow that is decreasing with radius, and falling below half at  $\sim 1$  kpc, but with overall similar molecular and atomic outflowing mass. Less observationally constrained is a path to very hot  $\gtrsim 10^7$  K dilute gas that is not easily detected in common X-ray data ([Strickland & Heckman 2009](#)). However, only for the fastest observed molecular outflows could the post-shock temperature  $T = \frac{3}{16} \frac{\mu v_s^2}{k}$  reach beyond  $10^7$  K, and potentially high local densities may favor re-cooling. Finally, some outflowing molecular gas could be converted to stars ([Zubovas & King 2014](#); [El-Badry et al. 2016](#); [Wang & Loeb 2018](#)), but observational confirmation is still scarce ([Maiolino et al. 2017](#); [Gallagher et al.](#)

[2019](#); [Rodríguez del Pino et al. 2019](#)), not yet having reached unambiguous detection of young stars that are both moving at the outflow speed and trapping a significant fraction of the outflow rate. Very large masses of young stars above the molecular outflow mass would build up, and should be little obscured, since the outflows typically extend beyond the dusty galaxy centers. In summary, there seems to be no obvious storage of molecular outflow in other phases that would help to drastically shorten the observed molecular gas flow times.

A second argument for time-dependent AGN driving rests on the energy and momentum budget. For about 10% of their sample, [Fluetsch et al. \(2019\)](#) argue that outflow kinetic power and momentum cannot be driven by the combination of the currently available AGN luminosity and SFR – a past episode of higher AGN luminosity must have been present.

Finally and most directly, in a few spatially well resolved cases, distinct outflow episodes can be inferred from the morphology. NGC 2623 shows both a clear bipolar outflow with the blue lobe north of the nucleus, and a second “fossil” blue component at a greater offset to the southeast (see Sect. 3, Fig. A.3). Flow times for these two components differ by more than an order of magnitude: 1 Myr versus 13 Myr. Rapidly evolving near-nuclear conditions of this advanced merger or driving by different nuclei (but note the single component appearance in *K*-band, [Rossa et al. 2007](#)) might be responsible for the different orientation of these two flows. For IRAS 08572+3915, deep IRAM-NOEMA data detect a second spatially separated outflow cloud that is offset by  $\sim 6$  kpc ([Janssen 2016](#); [Herrera-Camus et al. 2020b](#)) in addition to the previously known outflow ([Cicone et al. 2014](#)). On a smaller scale, spatially resolved ALMA data of ESO 302–G030 ([Pereira-Santaella et al. 2016](#), their Fig. 6) show several clumps over a total scale of  $\sim 0.8$  kpc, with a rather well defined blue–red symmetry. Driving by individual supernovae may just be energetically possible in this case, but intermittent AGN emission appears equally plausible.

Accretion rate variations on many timescales are a characteristic of the AGN phenomenon. Beyond the timescales of direct flux monitoring programs or archival light curve studies, their imprint on local AGN phenomena has been demonstrated in various ways. Light echoes of past activity can be seen in the form of (extended) narrow line regions that are too powerful for the currently observed AGN luminosity (e.g., [Lintott et al. 2009](#); [Keel et al. 2012](#); [Schirmer et al. 2013](#); see also [Sartori et al. 2018](#)). Conversely, AGNs with deficient NLR compared to their X-ray luminosity may trace the onset of activity ([Schawinski et al. 2015](#)). Perhaps most directly, X-ray light echoes trace accretion variations of the (much less luminous) activity in our Galactic Center (e.g., [Clavel et al. 2013](#)). Recent simulations start to reproduce the observed “flickering” AGN accretion variations (e.g., [Novak et al. 2011](#); [Gabor & Bournaud 2013](#)). At high redshift, the smoothing of accretion rate variations provided by the observed outflow can explain the higher incidence of AGN-driven nuclear outflows compared to direct AGN indicators such as X-rays ([Genzel et al. 2014](#); [Förster Schreiber et al. 2019](#)).

Molecular outflow statistics in a complete mass-selected sample could provide further insights into the timescales of AGN driving. Limitations arise from the additional role of star-formation-driven outflows that can be hard to discern from fossil AGN-driven outflows, and from the diverse target selection of outflow studies in the current literature. Tentatively, both the identification of multiple episodes in a few targets and the large outflow incidence (13 out of 17) for the [Lutz et al. \(2016\)](#)

high-surface-brightness sample argue that the  $\sim 1$  Myr outflow events may be episodes in a longer outflow history. However, the latter sample is a special population that favors a compact dense circumnuclear ISM, and the low incidence of OH outflows in the Seyfert sample of Stone et al. (2016) clearly shows that its statistics cannot be transferred to AGNs in general. In any case, “depletion times” needed to expel the entire molecular gas mass of a galaxy have to be viewed with great caution if one assumes that the current outflow rate persists unchanged over a very long time.

## 6. Conclusions

We have presented new CO(1-0) interferometry from ALMA and NOEMA, probing molecular outflows in 13 local galaxies with high far-infrared surface brightness, and combined our results with literature data on local universe molecular outflows. Our main findings are:

- (1) Galaxies with high far-infrared surface brightness  $\log(\Sigma_{\text{FIR}}) \geq 11.75 L_{\odot} \text{ kpc}^{-2}$  show the expected high incidence of molecular outflows. Suitable CO data is now available for 17 of the 18 sources in the Lutz et al. (2016) sample, and 13 of these exhibit molecular outflows. Several of these nearby targets, in particular III Zw 035, are well-suited to spatially resolved studies.
- (2) Partly constrained by observed line ratios and by spatial structure of the best cases, we adopt the following constants for the conversion of CO observables to outflow physical properties:  $\alpha_{\text{CO}(1-0)} = 0.8 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$  and  $C = 1$  in  $\dot{M}_{\text{out}} = C \frac{\dot{M}_{\text{out}} v_{\text{out}}}{R_{\text{out}}}$ . The high  $R_{31} = 2.1$  in IRAS 13120–5453 clearly indicates that this outflow is optically thin in CO. This suggests that  $\alpha_{\text{CO}(1-0)}$  may vary between outflows, but the good agreement with independent OH-based results globally supports the adopted CO conversion factor for outflows. We adopt conservative definitions of outflow flux and outflow velocity that focus on the line wings. While this may miss slow outflow components, which are likely to fall back, it improves robustness against confusion of outflow with undisturbed host gas.
- (3) We compare our results with outflow properties that are based on *Herschel* far-infrared spectra of OH absorption, with special focus on the 12 ULIRGs with detailed OH models by González-Alfonso et al. (2017). The good agreement in derived properties is an important cross-validation of the two main methods for characterizing molecular outflows. Modest remaining differences may relate to different but overlapping regions sampled by the two methods: smaller for the OH absorption versus larger for CO emission.
- (4) Outflow properties correlate better with AGN luminosity and with bolometric luminosity than with far-infrared surface brightness. The outflows with the greatest mass and mass outflow rate prefer systems with current luminous AGN activity, suggesting an important role of AGN driving, but significant outflows in some nonAGN systems must relate to star formation or to AGN activity in the recent past.
- (5) We report scaling relations of outflow properties with bolometric luminosity (Eqs. (4)–(7)), but note that these might still be influenced by selection effects.
- (6) Short  $\sim 10^6$  yr flow times and some sources with spatially resolved multiple outflow episodes support a role of intermittent or episodic driving, likely by AGNs.

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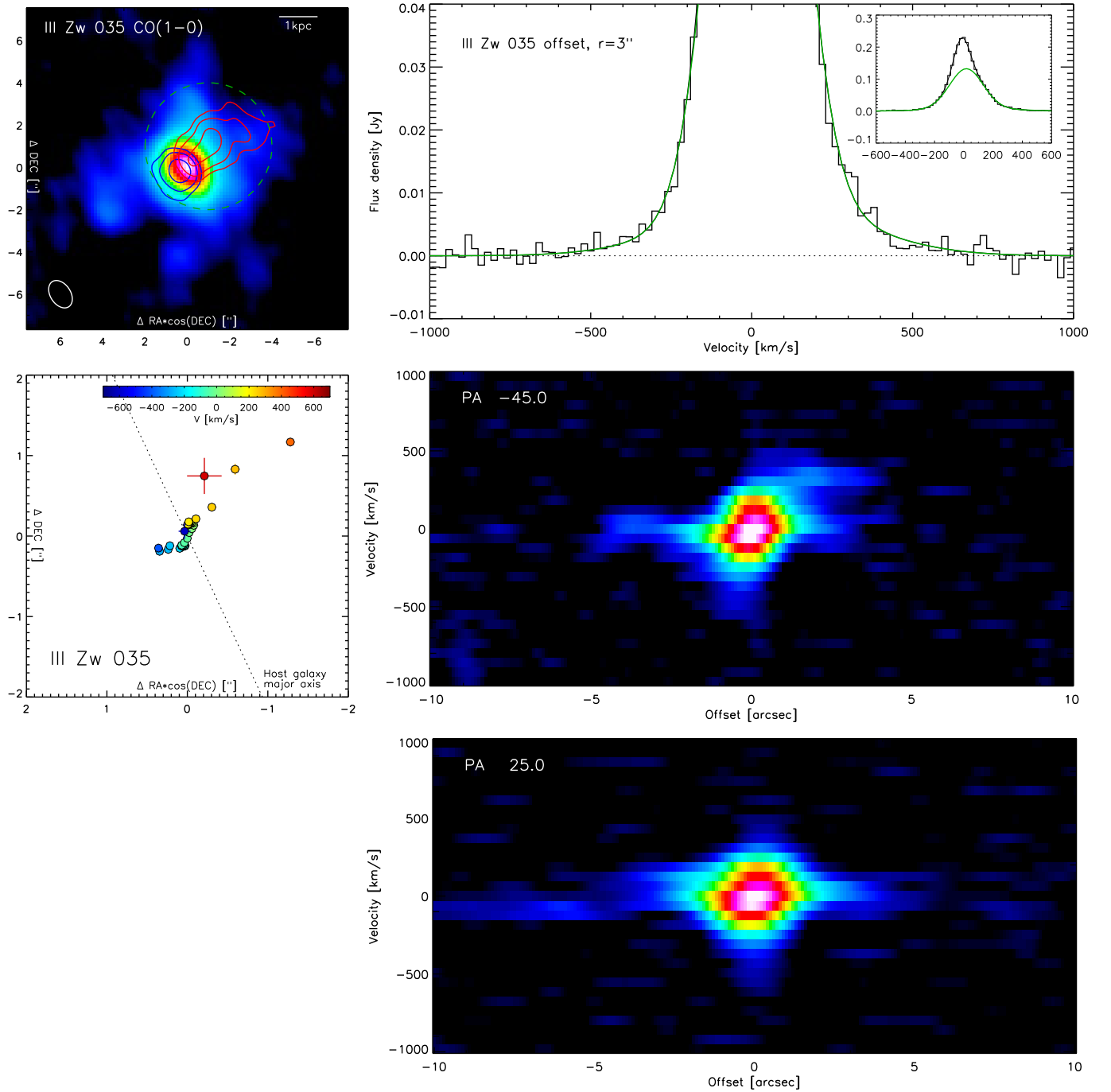
## References

- Aalto, S., García-Burillo, S., Müller, S., et al. 2012a, *A&A*, **537**, A44  
Aalto, S., Müller, S., Sakamoto, K., et al. 2012b, *A&A*, **546**, A68  
Aalto, S., García-Burillo, S., Müller, S., et al. 2015, *A&A*, **574**, A85  
Aalto, S., Costagliola, F., Müller, S., et al. 2016, *A&A*, **590**, A73  
Aladro, R., Martín, S., Riquelme, D., et al. 2015, *A&A*, **579**, A101  
Alatalo, K., Blitz, L., Young, L. M., et al. 2011, *ApJ*, **735**, 88  
Alonso-Herrero, A., Pereira-Santaella, M., Rieke, G. H., & Rigopoulou, D. 2012, *ApJ*, **744**, 2  
Alonso-Herrero, A., García-Burillo, S., Pereira-Santaella, M., et al. 2019, *A&A*, **628**, A65  
Baker, A. J. 2000, PhD Thesis, California Institute of Technology, USA  
Barcos-Muñoz, L., Leroy, A. K., Evans, A. S., et al. 2017, *ApJ*, **843**, 117  
Barcos-Muñoz, L., Aalto, S., Thompson, T. A., et al. 2018, *ApJ*, **853**, L28  
Bischetti, M., Piconcelli, E., Feruglio, C., et al. 2019, *A&A*, **628**, A118  
Bolatto, A. D., Wolfire, M., & Leroy, A. K. 2013a, *ARA&A*, **51**, 207  
Bolatto, A. D., Warren, S. R., Leroy, A. K., et al. 2013b, *Nature*, **499**, 450  
Bothwell, M. S., Smail, I., Chapman, S. C., et al. 2013, *MNRAS*, **429**, 3047  
Brightman, M., & Nandra, K. 2011, *MNRAS*, **414**, 3084  
Chabrier, G. 2003, *PASP*, **115**, 763  
Cicone, C., Feruglio, C., Maiolino, R., et al. 2012, *A&A*, **543**, A99  
Cicone, C., Maiolino, R., Sturm, E., et al. 2014, *A&A*, **562**, A21  
Cicone, A., Severgnini, P., Papadopoulos, P., et al. 2018, *ApJ*, **863**, 143  
Clavel, M., Terrier, R., Goldwurm, A., et al. 2013, *A&A*, **558**, A32  
Combes, F., García-Burillo, S., Casasola, V., et al. 2013, *A&A*, **558**, A124  
Combes, F., García-Burillo, S., Casasola, V., et al. 2014, *A&A*, **565**, A97  
Combes, F., García-Burillo, S., Audibert, A., et al. 2019, *A&A*, **623**, A79  
Contursi, A., Poglitsch, A., Gracia Carpio, J., et al. 2013, *A&A*, **549**, A118  
Costagliola, F., Sakamoto, K., Müller, S., et al. 2015, *A&A*, **582**, A91  
Cox, T. J., Dutta, S. N., Di Matteo, T., et al. 2006, *ApJ*, **650**, 791  
Dasyra, K. M., & Combes, F. 2012, *A&A*, **541**, L7  
Dasyra, K. M., Combes, F., Novak, G. S., et al. 2014, *A&A*, **565**, A46  
Dasyra, K. M., Combes, F., Oosterloo, T., et al. 2016, *A&A*, **595**, L7  
Davies, R. L., Förster Schreiber, N. M., Übler, H., et al. 2019, *ApJ*, **873**, 122  
El-Badry, K., Wetzel, A., Geha, M., et al. 2016, *ApJ*, **820**, 131  
Evans, A. S., Vavilkin, T., Pizagno, J., et al. 2008, *ApJ*, **675**, L69  
Fabian, A. C. 2012, *ARA&A*, **50**, 455  
Falstad, N., González-Alfonso, E., Aalto, S., et al. 2015, *A&A*, **580**, A52  
Falstad, N., Aalto, S., Mangum, J. G., et al. 2018, *A&A*, **609**, A75  
Falstad, N., Hallqvist, F., Aalto, S., et al. 2019, *A&A*, **623**, A29  
Faucher-Giguère, C.-A., & Quataert, E. 2012, *MNRAS*, **425**, 605  
Feruglio, C., Maiolino, R., Piconcelli, E., et al. 2010, *A&A*, **518**, L155  
Feruglio, C., Fiore, F., Piconcelli, E., et al. 2013, *A&A*, **558**, A87  
Feruglio, C., Fiore, F., Carniani, S., et al. 2015, *A&A*, **583**, A99  
Fiore, F., Feruglio, C., Shankar, F., et al. 2017, *A&A*, **601**, A143  
Fischer, J., Sturm, E., González-Alfonso, E., et al. 2010, *A&A*, **518**, L41  
Fluetsch, A., Maiolino, R., Carniani, S., et al. 2019, *MNRAS*, **483**, 4586  
Förster Schreiber, N. M., Übler, H., Davies, R. L., et al. 2019, *ApJ*, **875**, 21  
Gabor, J. M., & Bournaud, F. 2013, *MNRAS*, **434**, 606  
Gallagher, R., Maiolino, R., Belfiore, F., et al. 2019, *MNRAS*, **485**, 3409  
Gallimore, J. F., Elitzur, M., Maiolino, R., et al. 2016, *ApJ*, **829**, L7  
García-Burillo, S., Combes, F., Usero, A., et al. 2014, *A&A*, **567**, A125  
García-Burillo, S., Combes, F., Usero, A., et al. 2015, *A&A*, **580**, A35  
Genzel, R., Förster Schreiber, N. M., Rosario, D., et al. 2014, *ApJ*, **796**, 7  
Giroletti, M., Panessa, F., Longinotti, A. L., et al. 2017, *A&A*, **600**, A87  
González-Alfonso, E., Fischer, J., Graciá-Carpio, J., et al. 2012, *A&A*, **541**, A4  
González-Martín, O., Masegosa, J., Márquez, I., et al. 2015, *A&A*, **578**, A74  
González-Alfonso, E., Fischer, J., Spoon, H. W. W., et al. 2017, *ApJ*, **836**, 11  
Gowardhan, A., Spoon, H., Riechers, D., et al. 2018, *ApJ*, **859**, 35

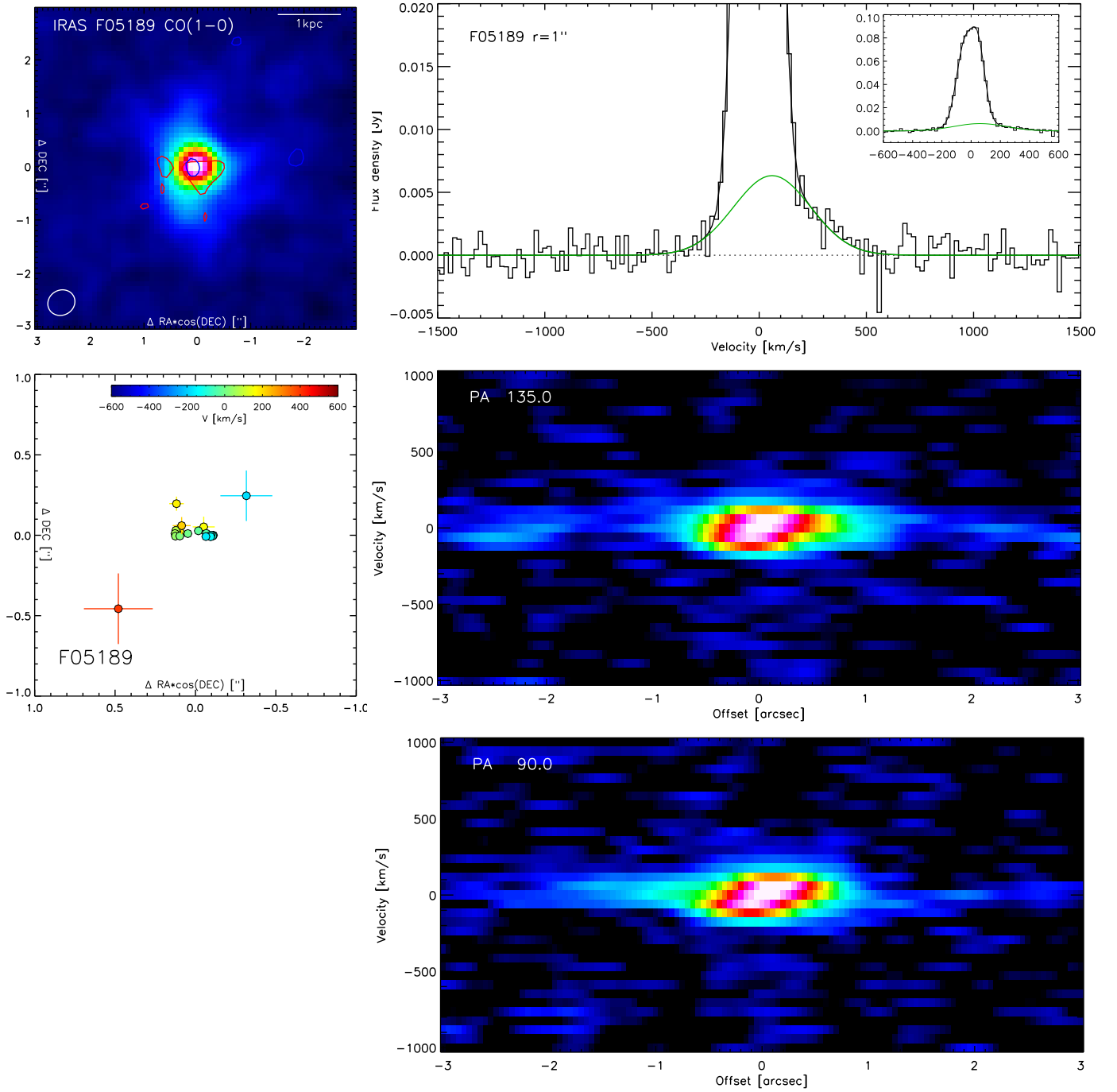


- Guilloteau, S., & Lucas, R. 2000, in *Imaging at Radio through Submillimeter Wavelengths*, eds. J. G. Mangum, & S. J. E. Radford, *ASP Conf. Ser.*, **217**, 299
- Hayward, C. C., Torrey, P., Springel, V., Hernquist, L., & Vogelsberger, M. 2014, *MNRAS*, **442**, 1992
- Heckman, T. M. 2002, in *Extragalactic Gas at Low Redshift*, eds. J. S. Mulchaey, & J. T. Stocke, *ASP Conf. Ser.*, **254**, 292
- Herrera-Camus, R., Tacconi, L., Genzel, R., et al. 2019, *ApJ*, **871**, 37
- Herrera-Camus, R., Sturm, E., Graciá-Carpio, J., et al. 2020a, *A&A*, **633**, L4
- Herrera-Camus, R., Janssen, A., Sturm, E., et al. 2020b, *A&A*, in press, <https://doi.org/10.1051/0004-6361/201936434>
- Hopkins, P. F., Cox, T. J., Kereš, D., & Hernquist, L. 2008, *ApJS*, **175**, 390
- Hopkins, P. F., Cox, T. J., Hernquist, L., et al. 2013, *MNRAS*, **430**, 1901
- Iwasawa, K., Sanders, D. B., Teng, S. H., et al. 2011, *A&A*, **529**, A106
- Janssen, A. 2016, PhD Thesis, Ludwig-Maximilians-Universität München, Germany
- Janssen, A. W., Christopher, N., Sturm, E., et al. 2016, *ApJ*, **822**, 43
- Keel, W. C., Chojnowski, S. D., Bennert, V. N., et al. 2012, *MNRAS*, **420**, 878
- Kelly, B. C. 2007, *ApJ*, **665**, 1489
- Kennicutt, Jr., R. C. 1998, *ARA&A*, **36**, 189
- Kim, D.-C., Evans, A. S., Vavilkin, T., et al. 2013, *ApJ*, **768**, 102
- King, A., & Pounds, K. 2015, *ARA&A*, **53**, 115
- Laine, S., Kenney, J. D. P., Yun, M. S., & Gottesman, S. T. 1999, *ApJ*, **511**, 709
- Leon, S., Eckart, A., Laine, S., et al. 2007, *A&A*, **473**, 747
- Leroy, A. K., Walter, F., Martini, P., et al. 2015a, *ApJ*, **814**, 83
- Leroy, A. K., Walter, F., Decarli, R., et al. 2015b, *ApJ*, **811**, 15
- Lintott, C. J., Schawinski, K., Keel, W., et al. 2009, *MNRAS*, **399**, 129
- Longinotti, A. L., Vega, O., Krongold, Y., et al. 2018, *ApJ*, **867**, L11
- Lutz, D., Berta, S., Contursi, A., et al. 2016, *A&A*, **591**, A136
- Lutz, D., Shimizu, T., Davies, R. I., et al. 2018, *A&A*, **609**, A9
- Maiolino, R., Russell, H. R., Fabian, A. C., et al. 2017, *Nature*, **544**, 202
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in *Astronomical Data Analysis Software and Systems XVI*, eds. R. A. Shaw, F. Hill, & D. J. Bell, *ASP Conf. Ser.*, **376**, 127
- Meier, D. S., Walter, F., Bolatto, A. D., et al. 2015, *ApJ*, **801**, 63
- Mihos, J. C., & Hernquist, L. 1996, *ApJ*, **464**, 641
- Murray, N., Quataert, E., & Thompson, T. A. 2005, *ApJ*, **618**, 569
- Nardini, E., Risaliti, G., Watabe, Y., Salvati, M., & Sani, E. 2010, *MNRAS*, **405**, 2505
- Newman, S. F., Genzel, R., Förster-Schreiber, N. M., et al. 2012, *ApJ*, **761**, 43
- Novak, G. S., Ostriker, J. P., & Ciotti, L. 2011, *ApJ*, **737**, 26
- Pereira-Santaella, M., Colina, L., García-Burillo, S., et al. 2016, *A&A*, **594**, A81
- Pereira-Santaella, M., Colina, L., García-Burillo, S., et al. 2018, *A&A*, **616**, A171
- Privon, G. C., Aalto, S., Falstad, N., et al. 2017, *ApJ*, **835**, 213
- Querejeta, M., Schinnerer, E., García-Burillo, S., et al. 2016, *A&A*, **593**, A118
- Ramakrishnan, V., Nagar, N. M., Finlez, C., et al. 2019, *MNRAS*, **487**, 444
- Ricci, C., Trakhtenbrot, B., Koss, M. J., et al. 2017, *ApJS*, **233**, 17
- Rodríguez del Pino, B., Arribas, S., Piqueras López, J., Villar-Martín, M., & Colina, L. 2019, *MNRAS*, **486**, 344
- Roelfsema, P. R., Shibai, H., Armus, L., et al. 2018, *PASA*, **35**, e030
- Rossa, J., Laine, S., van der Marel, R. P., et al. 2007, *AJ*, **134**, 2124
- Rupke, D. S. N., & Veilleux, S. 2013, *ApJ*, **768**, 75
- Rupke, D. S., Veilleux, S., & Sanders, D. B. 2005, *ApJS*, **160**, 115
- Rupke, D. S. N., Gültekin, K., & Veilleux, S. 2017, *ApJ*, **850**, 40
- Sakamoto, K., Aalto, S., Costagliola, F., et al. 2013, *ApJ*, **764**, 42
- Sakamoto, K., Aalto, S., Combes, F., Evans, A., & Peck, A. 2014, *ApJ*, **797**, 90
- Salak, D., Nakai, N., Hatakeyama, T., & Miyamoto, Y. 2016, *ApJ*, **823**, 68
- Sartori, L. F., Schawinski, K., Trakhtenbrot, B., et al. 2018, *MNRAS*, **476**, L34
- Schawinski, K., Koss, M., Berney, S., & Sartori, L. F. 2015, *MNRAS*, **451**, 2517
- Schirmer, M., Diaz, R., Holmberg, K., Levenson, N. A., & Winge, C. 2013, *ApJ*, **763**, 60
- Sempere, M. J., Combes, F., & Casoli, F. 1995, *A&A*, **299**, 371
- Slater, R., Nagar, N. M., Schnorr-Müller, A., et al. 2019, *A&A*, **621**, A83
- Sliwa, K., Wilson, C. D., Aalto, S., & Privon, G. C. 2017, *ApJ*, **840**, L11
- Spilker, J. S., Aravena, M., Béthermin, M., et al. 2018, *Science*, **361**, 1016
- Spoon, H. W. W., Farrah, D., Leboutteiller, V., et al. 2013, *ApJ*, **775**, 127
- Stone, M., Veilleux, S., Meléndez, M., et al. 2016, *ApJ*, **826**, 111
- Stone, M., Veilleux, S., González-Alfonso, E., Spoon, H., & Sturm, E. 2018, *ApJ*, **853**, 132
- Strickland, D. K., & Heckman, T. M. 2009, *ApJ*, **697**, 2030
- Sturm, E., González-Alfonso, E., Veilleux, S., et al. 2011, *ApJ*, **733**, L16
- Sun, A.-L., Greene, J. E., Zakamska, N. L., & Nesvadba, N. P. H. 2014, *ApJ*, **790**, 160
- Tacconi, L. J., Genzel, R., Saintonge, A., et al. 2018, *ApJ*, **853**, 179
- Tsai, A.-L., Matsushita, S., Nakanishi, K., et al. 2009, *PASJ*, **61**, 237
- Tsai, A.-L., Matsushita, S., Kong, A. K. H., Matsumoto, H., & Kohno, K. 2012, *ApJ*, **752**, 38
- Veilleux, S., Cecil, G., & Bland-Hawthorn, J. 2005, *ARA&A*, **43**, 769
- Veilleux, S., Meléndez, M., Sturm, E., et al. 2013, *ApJ*, **776**, 27
- Veilleux, S., Bolatto, A., Tombesi, F., et al. 2017, *ApJ*, **843**, 18
- Walter, F., Weiss, A., & Scoville, N. 2002, *ApJ*, **580**, L21
- Walter, F., Bolatto, A. D., Leroy, A. K., et al. 2017, *ApJ*, **835**, 265
- Wang, X., & Loeb, A. 2018, *New Astron.*, **61**, 95
- Warmels, R., Biggs, A., Cortes, P. A., et al. 2018, *ALMA Technical Handbook*, *ALMA Doc 6.3*, ver 1.0
- Weiß, A., Walter, F., & Scoville, N. Z. 2005, *A&A*, **438**, 533
- Wuyts, S., Cox, T. J., Hayward, C. C., et al. 2010, *ApJ*, **722**, 1666
- Zschaechner, L. K., Bolatto, A. D., Walter, F., et al. 2018, *ApJ*, **867**, 111
- Zubovas, K., & King, A. 2012, *ApJ*, **745**, L34
- Zubovas, K., & King, A. R. 2014, *MNRAS*, **439**, 400
- Zubovas, K., & King, A. 2016, *MNRAS*, **462**, 4055

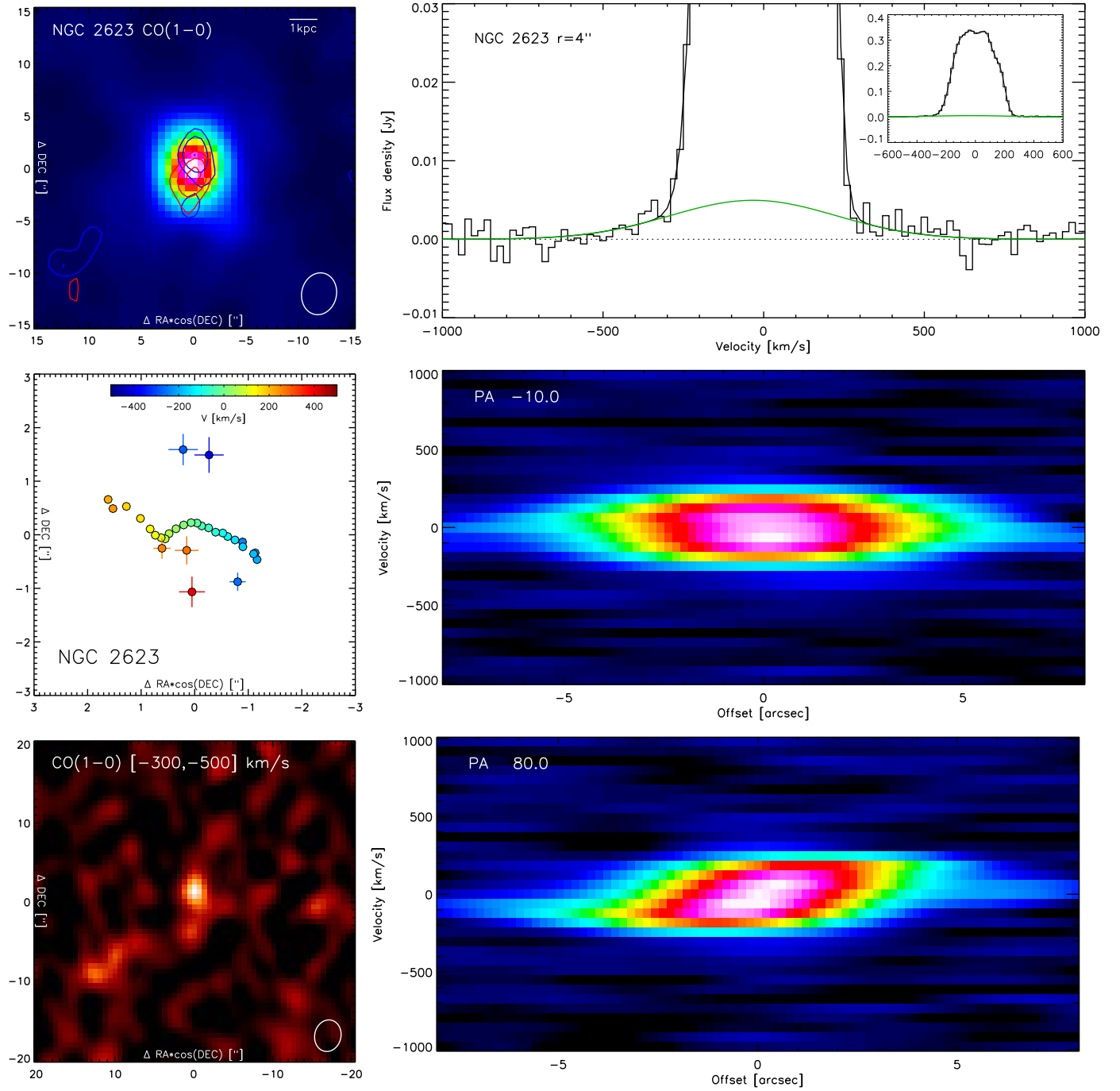
## Appendix A: Figures for individual sources



**Fig. A.1.** NOEMA CO(1-0) data for III Zw 035. *Top left:* moment-zero map. All spatial offsets are in arcseconds relative to the position of the 3 mm continuum (Table 2). North is to the top and east to the left. Overlaid contours are for outflow in the velocity range  $[-570, -270] \text{ km s}^{-1}$  (blue) and  $[270, 570] \text{ km s}^{-1}$  (red), with contours at  $[0.3, 0.6, 1.2] \text{ mJy beam}^{-1}$ . *Top right:* spectrum in a  $r = 3''$  aperture offset from the continuum nucleus by  $-1''$  in  $\text{RA} \times \cos(\text{Dec})$  and  $+1''$  in Dec (green dashed circle in *top left panel*). The spectrum is decomposed into three Gaussians, one for host and two for outflow (green). The black line includes all three Gaussian components. *Center left:* center positions and their errors for selected velocity channels. These are derived by fitting a Gaussian model to the UV data for the respective velocity channel. North is to the top and east to the left. The dotted line marks the galaxy major axis position angle  $\sim 25^\circ$  (Kim et al. 2013). *Center right:* position-velocity diagram along the direction of the strong red outflow at PA  $-45^\circ$ . *Bottom right:* position-velocity diagram along the galaxy major axis PA  $25^\circ$ . The fainter galaxy of the pair is located near offset  $-7''$  in this PV diagram but does not stand out clearly from the major emission of the galaxy. We detect a strong and asymmetric bipolar outflow.

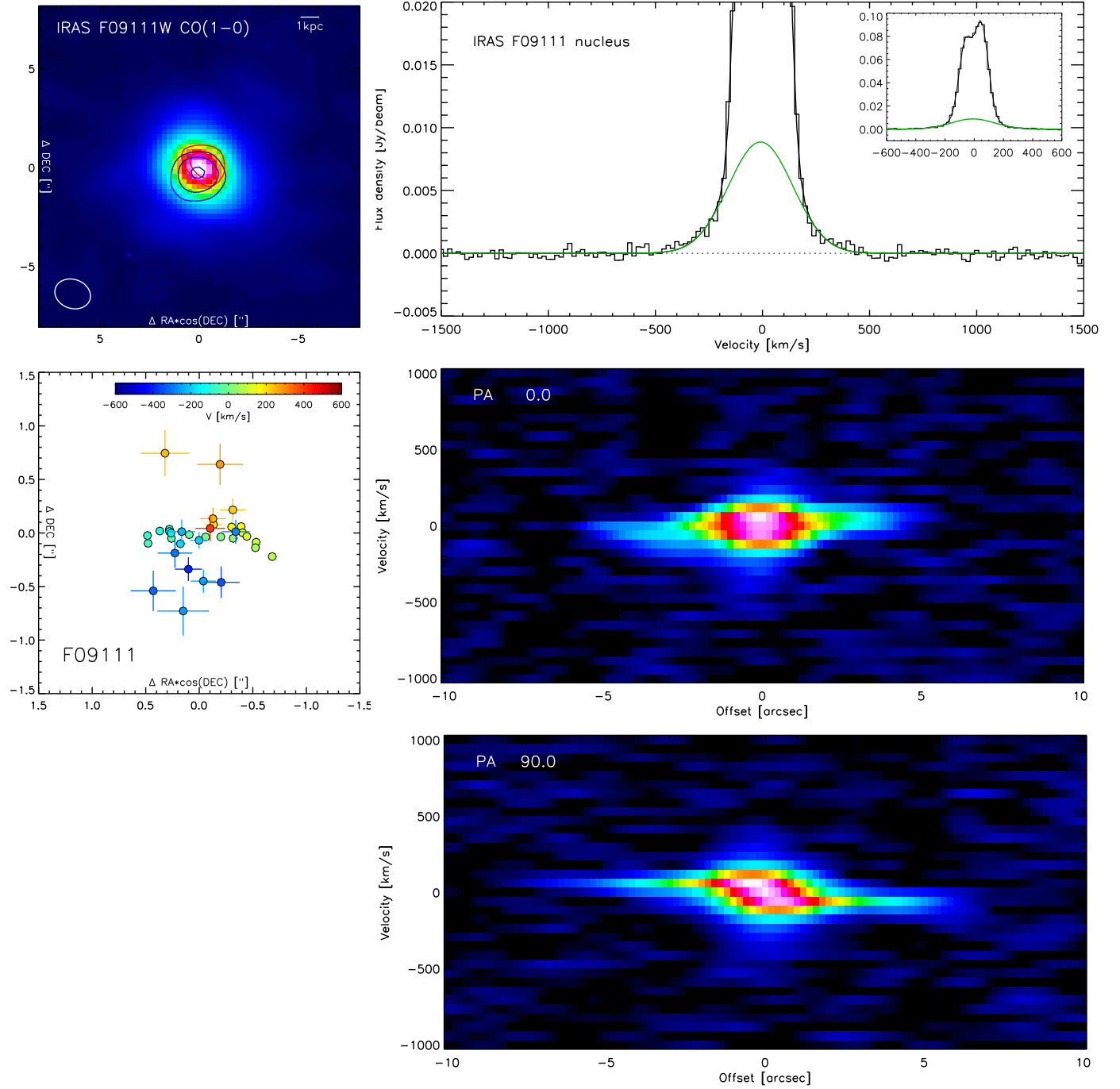


**Fig. A.2.** ALMA CO(1-0) data for IRAS F05189–2524. *Top left:* moment-zero map. Overlaid contours are for the outflow in the velocity range  $[210, 510] \text{ km s}^{-1}$  (red,  $0.3 \text{ mJy beam}^{-1}$ ) and the marginal emission at  $[-990, -270] \text{ km s}^{-1}$  (blue,  $0.2 \text{ mJy beam}^{-1}$ ). *Top right:* spectrum in a  $r = 1''$  aperture centered on the continuum nucleus. The spectrum is decomposed into three Gaussians, two for host and one for outflow (green). The black line includes all three Gaussian components. Outflow is clearly detected in the redshifted wing, and tentatively indicated for the blue wing. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual channels. *Center and bottom right:* position-velocity diagrams. PA +135° approximately follows the diffuse red outflow component towards the southeast or south, while PA +90° traces the velocity gradient of the line core.

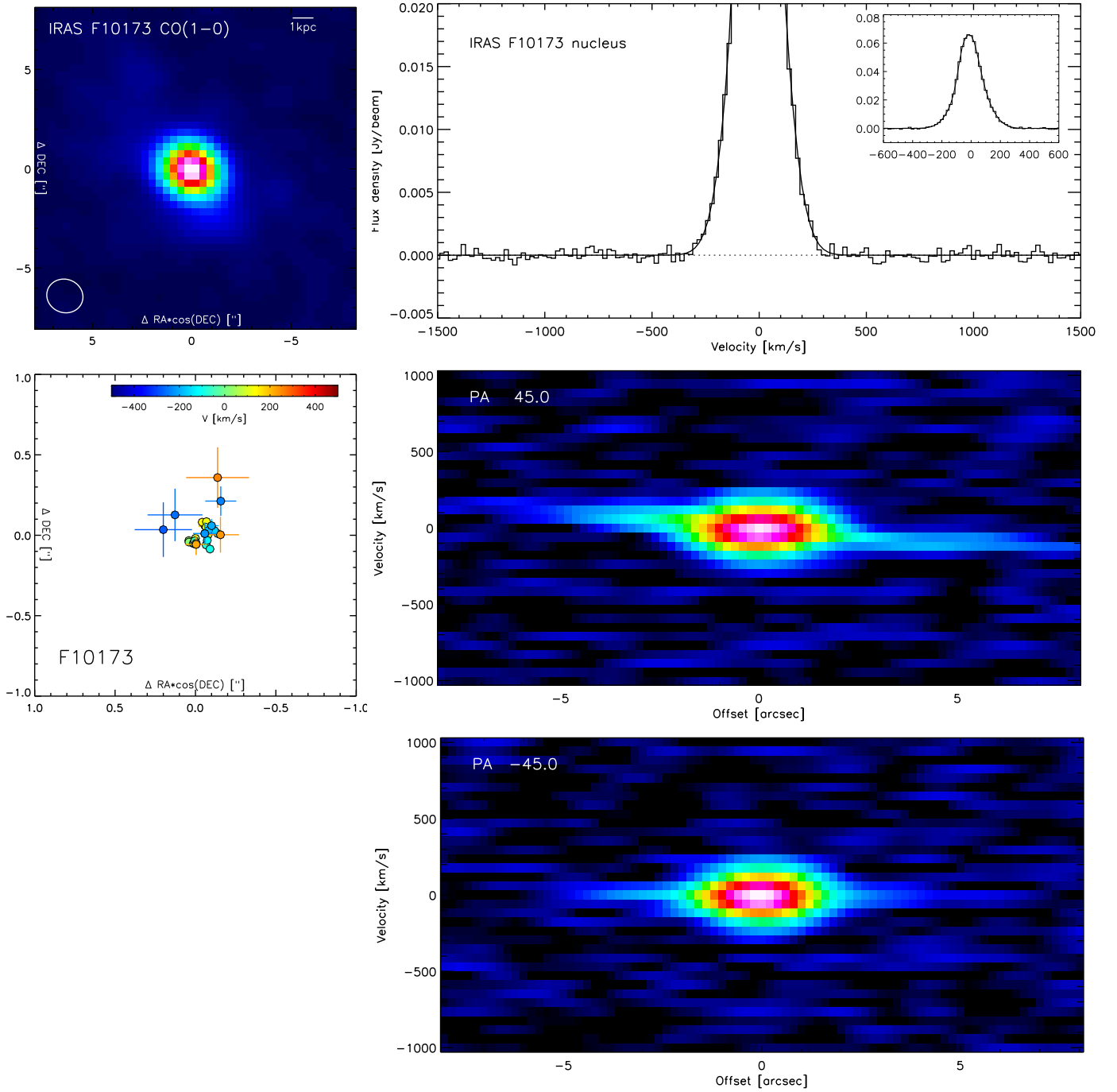


**Fig. A.3.** NOEMA CO(1-0) data for NGC 2623. *Top left:* moment-zero map. Overlaid contours are for the outflow in the velocity range  $[-500, -300]$  (blue) and  $[300, 500]$  km s $^{-1}$  (red), with contours at  $[0.7, 1.0, 1.5]$  mJy beam $^{-1}$ . *Top right:* spectrum in a  $r = 4''$  aperture centered on the continuum nucleus. The spectrum is decomposed into five Gaussians, four for the complex host line profile and one for outflow (green). The black line includes all five Gaussian components. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual channels. *Center and bottom right:* position-velocity diagrams. PA  $+80^\circ$  traces the velocity gradient of the line core. The orthogonal PA  $-10^\circ$  is also close to the orientation of the roughly N-S bipolar outflow seen in the *top left* and *center left* panels. *Bottom left:*  $[-500, -300]$  km s $^{-1}$  channel map showing the SE “fossil” outflow that is also indicated in the blue contours of the *top left* panel, in addition to the blue component of the main bipolar outflow.

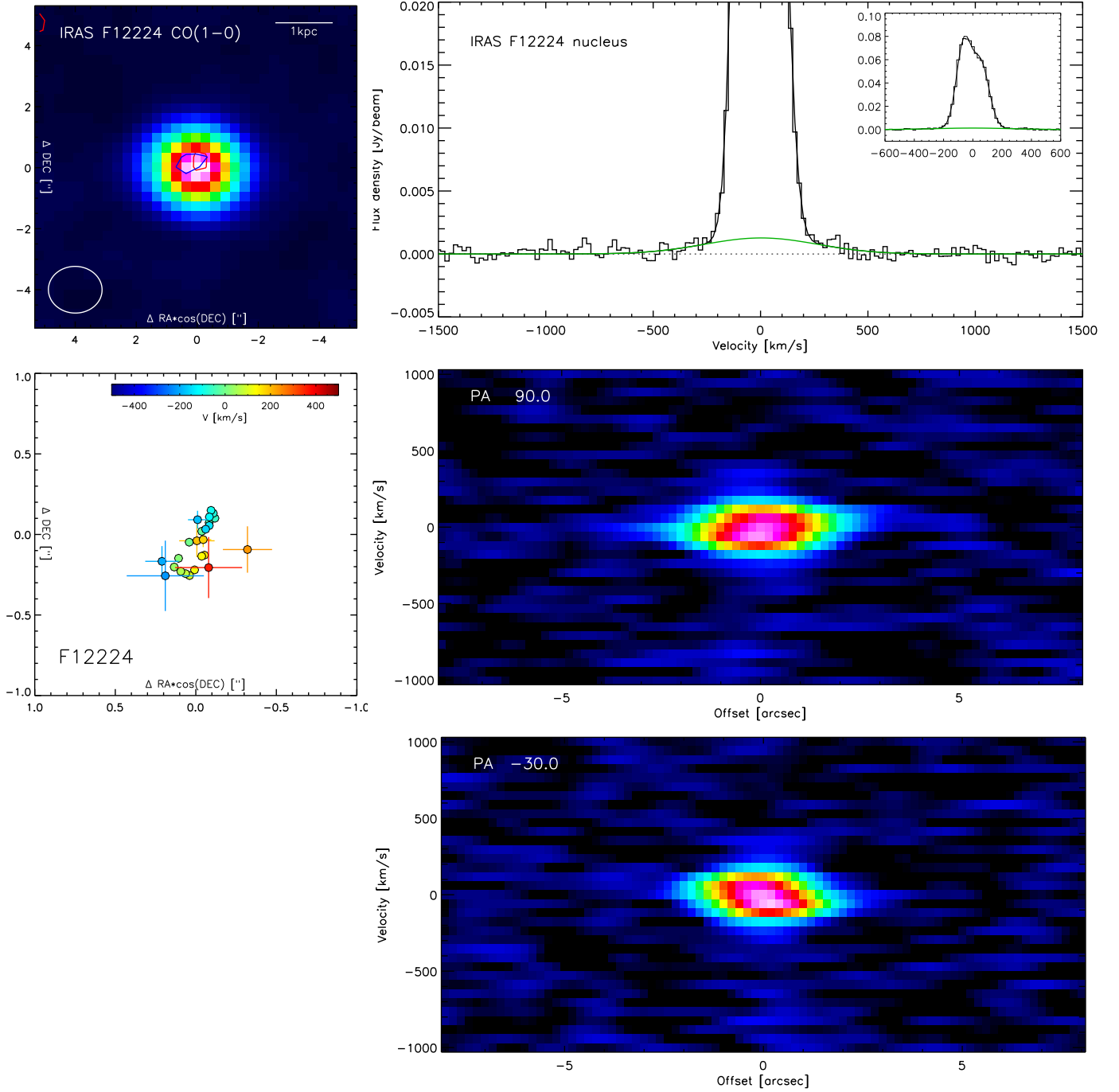




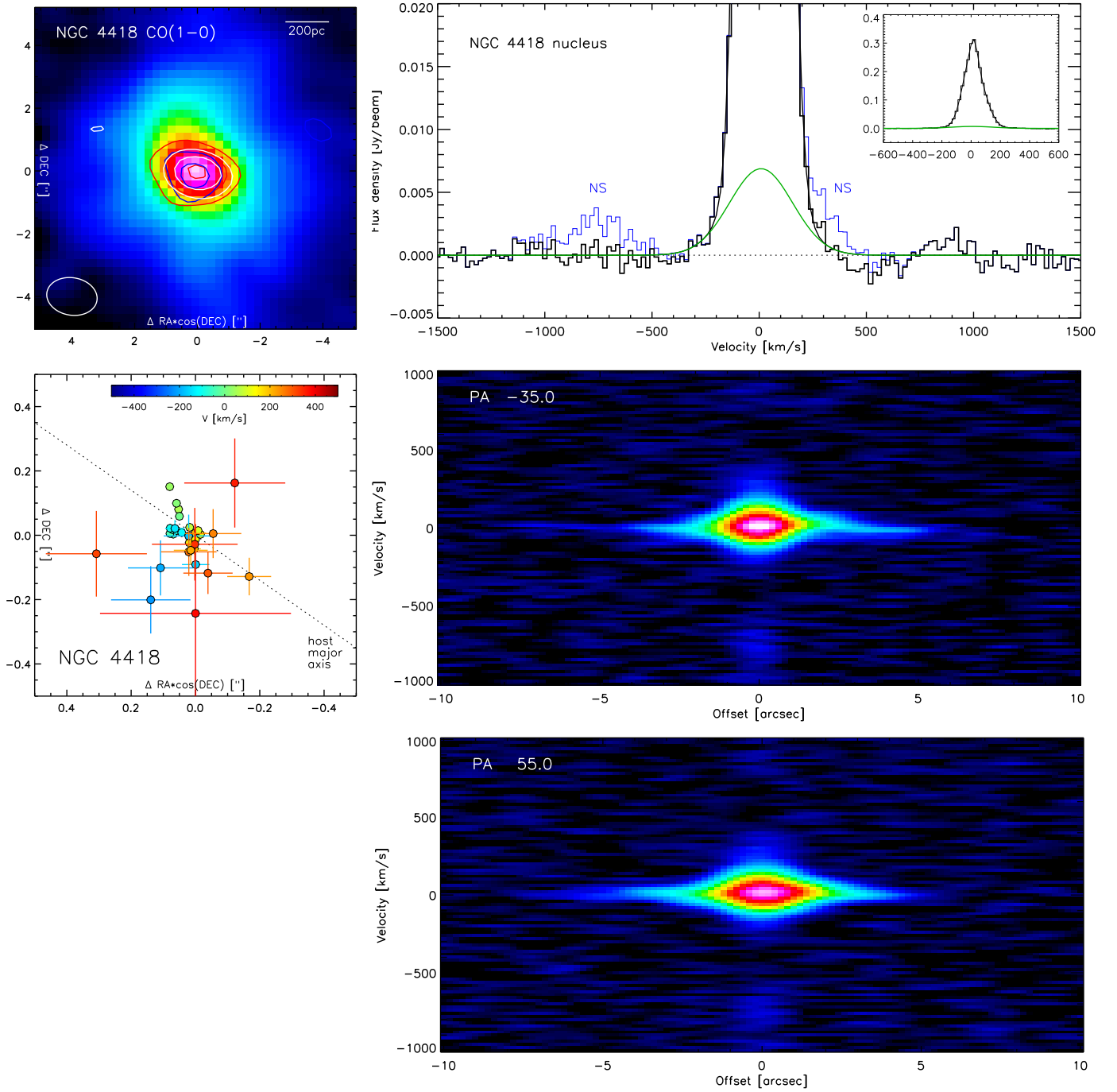
**Fig. A.4.** ALMA CO(1-0) data for IRAS F09111-1007W. *Top left:* moment-zero map. Overplotted are contours for outflow in the range  $[-510, -210]$  km s $^{-1}$  (blue) and  $[210, 510]$  km s $^{-1}$  (red), with contours at  $[0.35, 0.6, 1.1]$  mJy beam $^{-1}$ . *Top right:* spectrum centered on the continuum nucleus. The spectrum is decomposed into three Gaussians, two for the host line profile and one for outflow (green). The black line includes all three Gaussian components. Outflow wings are detected on both sides of the line core. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams. PA  $0^\circ$  is close to the bipolar outflow orientation, and orthogonal to the velocity gradient of the line core at PA  $\sim 90^\circ$ .



**Fig. A.5.** ALMA CO(1-0) data for IRAS F10173+0828. *Top left:* moment-zero map. *Top right:* spectrum centered on the continuum nucleus. The spectrum is fitted with three Gaussians which we all assign to the host. The black line includes all three Gaussian components. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams. In the absence of a clear outflow signature or of a clear velocity gradient in the line core, we simply plot two orthogonal directions of which PA +45° follows a velocity gradient on scales of several arcseconds. We conservatively adopt an upper limit for outflow in this source but cannot exclude that part of the line wings to  $\pm 300 \text{ km s}^{-1}$  are caused by outflow.

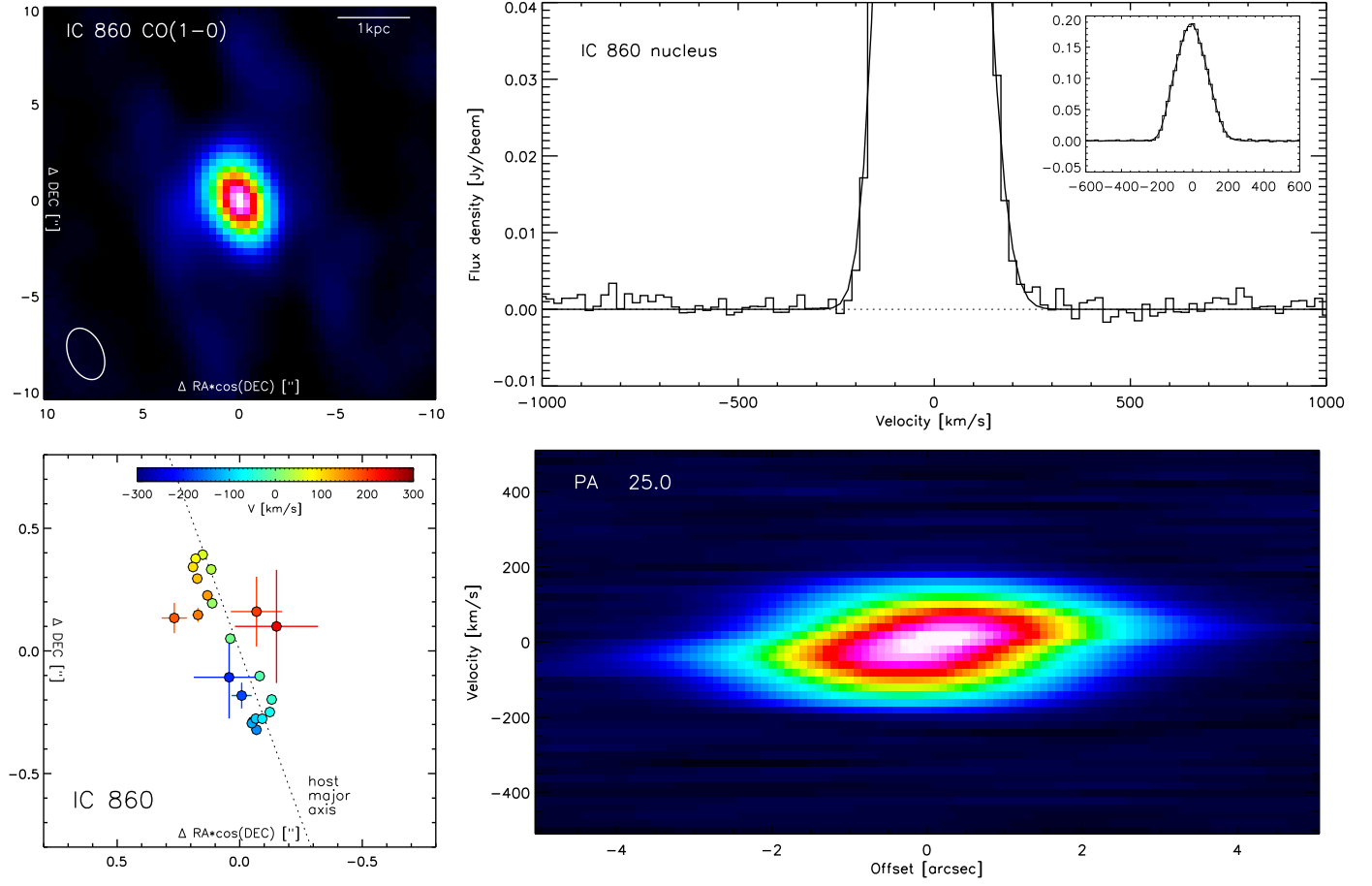


**Fig. A.6.** ALMA CO(1-0) data for IRAS F12224–0624. *Top left:* moment-zero map. Overplotted are contours ( $0.4 \text{ mJy} = 3.8\sigma$ ) for outflow in the range  $[-510, -270] \text{ km s}^{-1}$  (blue) and  $[270, 510] \text{ km s}^{-1}$  (red), with contours at  $0.4 \text{ mJy beam}^{-1}$ . *Top right:* spectrum centered on the continuum nucleus. The spectrum is decomposed into three Gaussians, two for the host line profile and one for outflow (green). The black line includes all three Gaussian components. Weak outflow wings are detected on the blue and red side of the line core. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams. PA  $+90^\circ$  follows the tentative E–W offset of blue and red outflow wings, while PA  $-30^\circ$  represents the velocity gradient in the line core.

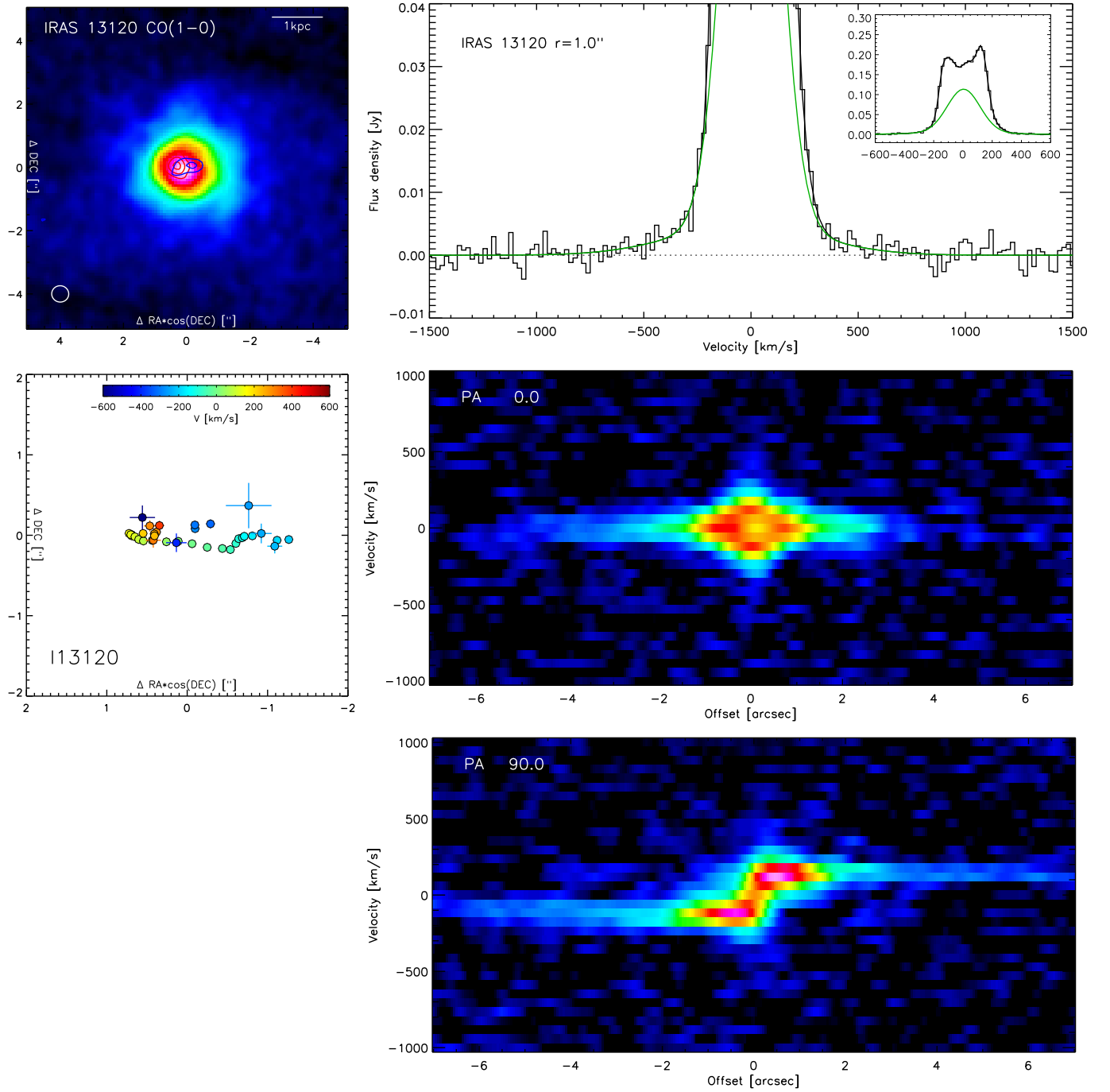


**Fig. A.7.** ALMA CO(1-0) data for NGC 4418. *Top left:* moment-zero map. Overplotted are contours for the ranges  $[-930, -570] \text{ km s}^{-1}$  (white, this range is dominated by emission of nitrogen sulfide),  $[-330, -190] \text{ km s}^{-1}$  (blue, blue wing of CO,  $[0.8, 1.6] \text{ mJy beam}^{-1}$ ) and  $[230, 510] \text{ km s}^{-1}$  (red, nitrogen sulfide NS superposed on red wing of CO,  $[0.6, 1.2, 2.4] \text{ mJy beam}^{-1}$ ). *Top right:* spectrum centered on the continuum nucleus. Before fitting, we have subtracted emission from nitrogen sulfide (blue) as described in the text, to obtain the black histogram CO(1-0) spectrum. CO outflow wings remain present in the line profile after subtraction of nitrogen sulfide. Nitrogen sulfide emission is still present in the other panels of this figure. The CO spectrum is decomposed into four Gaussians, three for the host line profile and one for outflow (green). The black line includes all four Gaussian components. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. The dotted line indicates the orientation of the major axis of the large scale disk of NGC 4418, according to 2MASS. *Center and bottom right:* position-velocity diagrams. PA  $-35^\circ$  is orthogonal to the host galaxy's major axis PA  $+55^\circ$ . No outflow orientation is firmly established given signal-to-noise and contamination by NS, but the lesser contaminated blue wing may prefer outflow orthogonal to the disk.

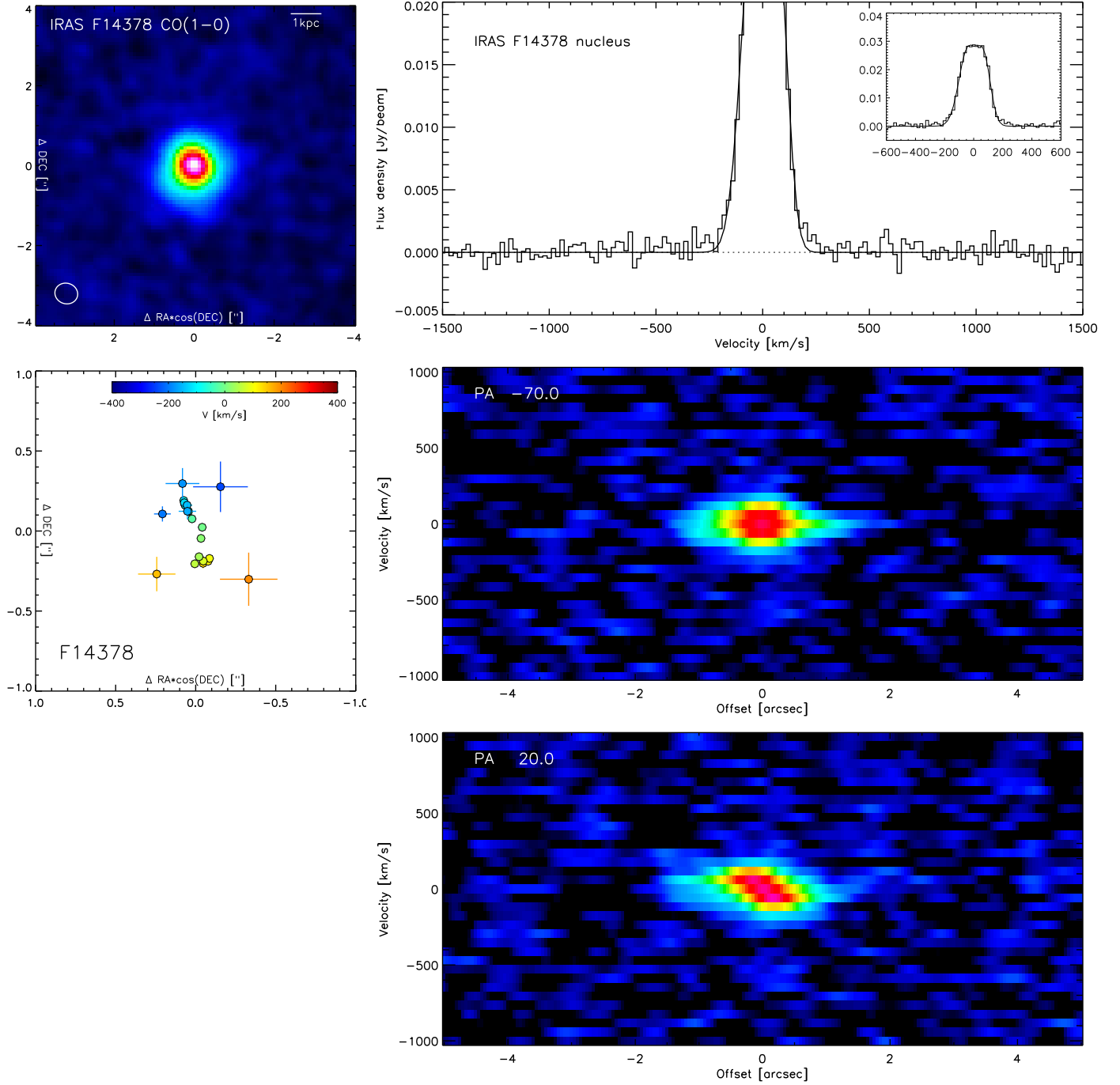




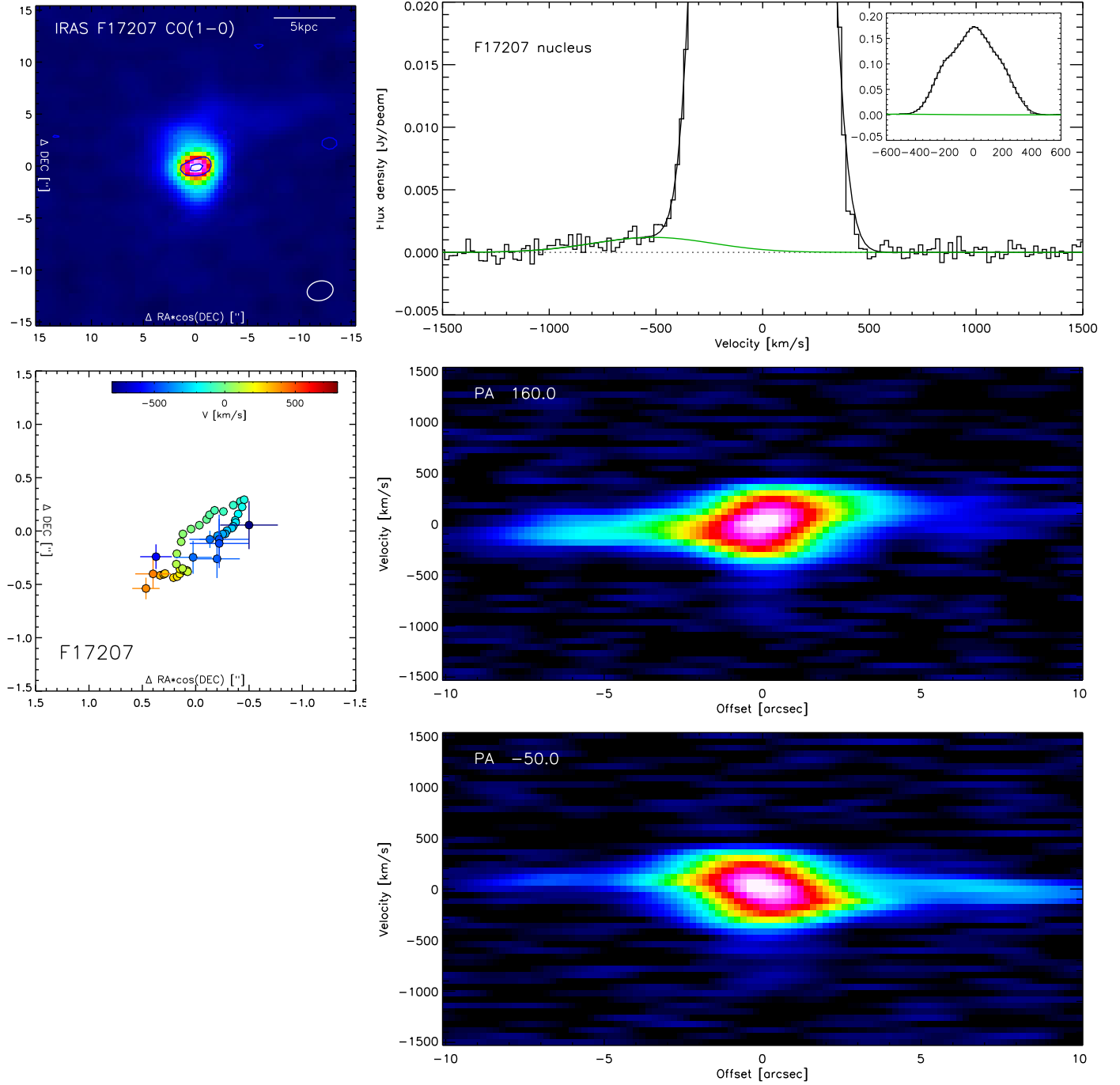
**Fig. A.8.** NOEMA CO(1-0) data for IC 860. *Top left:* moment-zero map. *Top right:* spectrum centered on the continuum nucleus. The host spectrum is fit with three Gaussians, their sum is overplotted in black. We conservatively assign an upper limit to a CO outflow. *Bottom left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. The dotted line indicates the 2MASS extended source catalog  $K$ -band major axis PA  $20^\circ$  of the stellar host. *Bottom right:* position-velocity diagram at PA  $+25^\circ$  deg, close to the velocity gradient of the line core and to the host major axis.



**Fig. A.9.** ALMA CO(1-0) data for IRAS 13120-5453. *Top left:* moment-zero map. Overplotted are contours for outflow in the range  $[-690, -270]$  km s $^{-1}$  (blue) and  $[270, 510]$  km s $^{-1}$  (red), with contours at  $[0.55, 0.9, 1.2]$  mJy beam $^{-1}$ . *Top right:* spectrum in a  $r = 1''$  aperture centered on the continuum nucleus. The spectrum is decomposed into six Gaussians, four for the double-horned host line profile and two for outflow (green). The black line includes all six Gaussian components. The source hosts a strong outflow component. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams, along the line core velocity gradient at PA  $\sim 90^\circ$  and orthogonal to it.

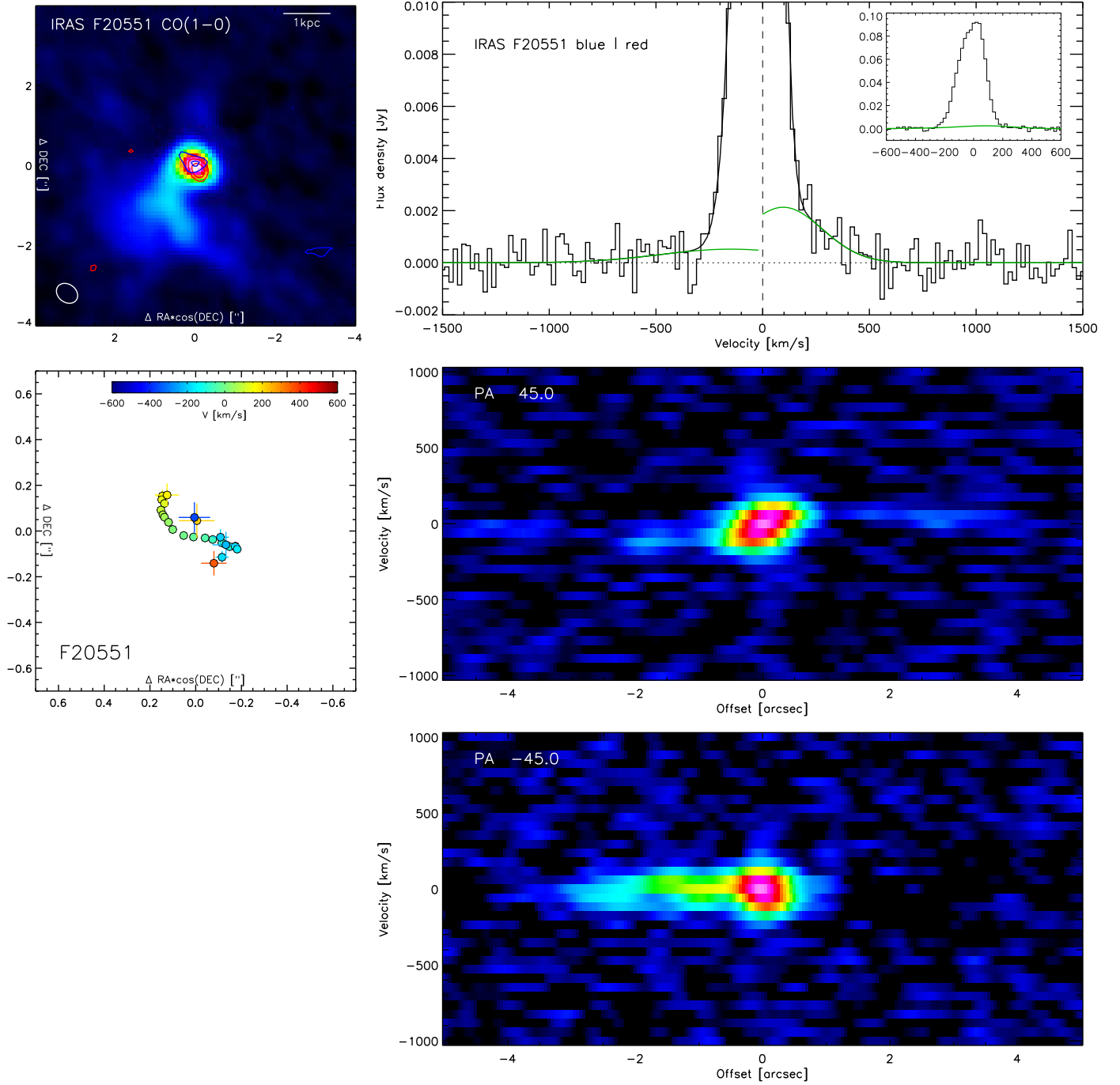


**Fig. A.10.** ALMA CO(1-0) data for IRAS F14378–3651. *Top left:* moment-zero map. *Top right:* spectrum centered on the continuum nucleus. The host spectrum is fit with two Gaussians, their sum is overplotted in black. We conservatively assign an upper limit to molecular outflow, since the nature of the weak line wings at  $\pm 250 \text{ km s}^{-1}$  is unclear. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams, along the line core velocity gradient at PA  $\sim 20^\circ$  and orthogonal to it.

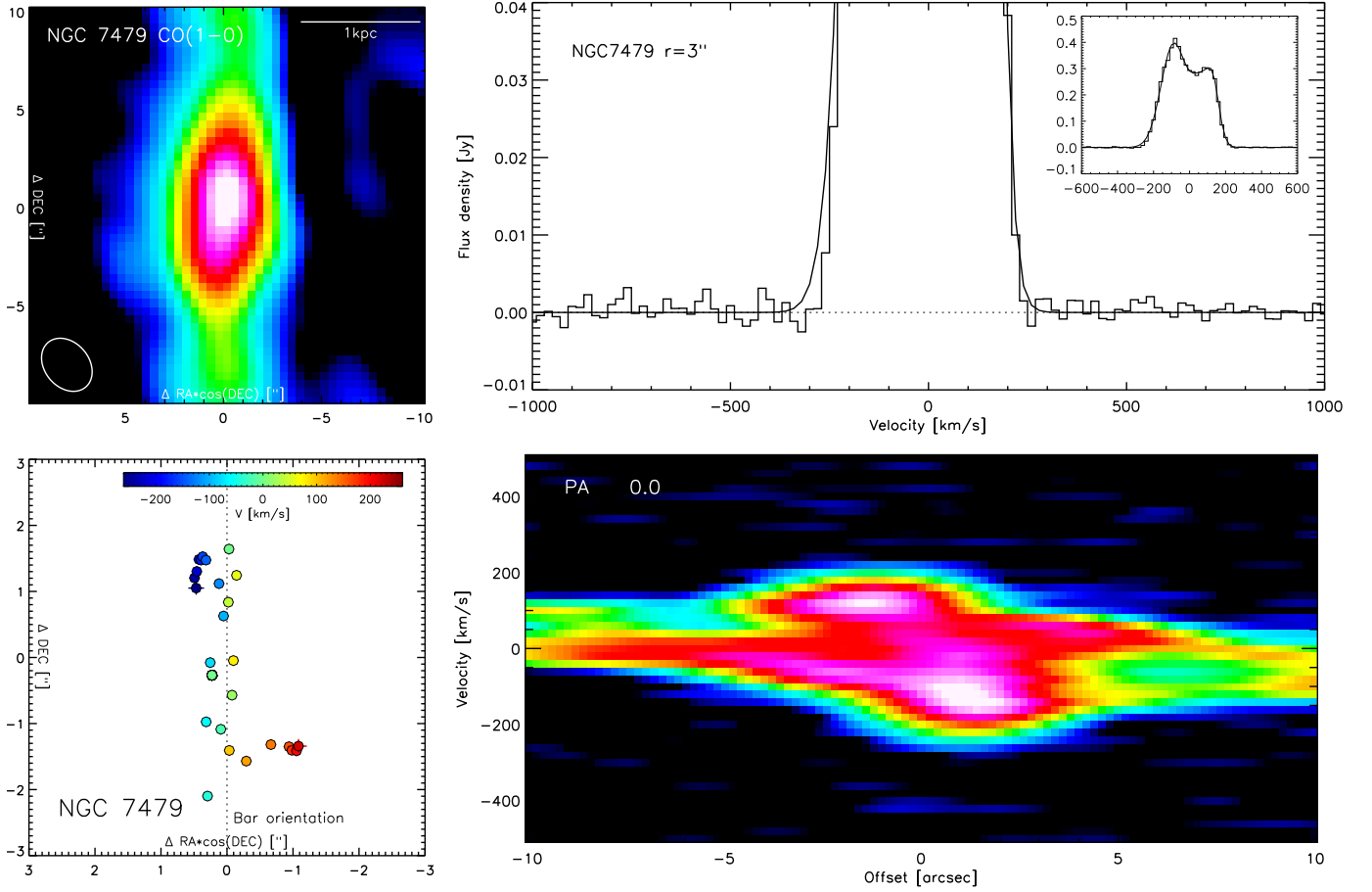


**Fig. A.11.** ALMA CO(1-0) data for IRAS F17207-0014. *Top left:* moment-zero map. Overplotted are contours for outflow in the range  $[-900, -500] \text{ km s}^{-1}$  (blue, contours at  $[0.6, 0.9] \text{ mJy beam}^{-1}$ ). *Top right:* nuclear spectrum. The spectrum is decomposed into four Gaussians, three for the host line profile and one for outflow (green). The black line includes all four Gaussian components. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams. Fast outflowing molecular gas is seen on the blue side only. PA  $+160^\circ$  is the orientation of the blue outflow component measured in higher-resolution CO(2-1) data by [García-Burillo et al. \(2015\)](#), our lower-resolution data are consistent with this by placing the blue outflow slightly south of the nucleus (see *center left panel*). PA  $-50^\circ$  is the approximate orientation of the central arcsecond-scale gas disk.





**Fig. A.12.** ALMA CO(1-0) data for IRAS F20551-4250. *Top left:* moment-zero map. Overplotted are contours for outflow in the range [210, 510] km s<sup>-1</sup> (red) and the tentative outflow in [-510, -210] km s<sup>-1</sup> (blue), with contours at [0.5, 0.7, 1.0] mJy beam<sup>-1</sup>. *Top right:* spectrum at the position of red outflow (right half) and blue outflow (left half). Each spectrum, including the part not shown here, is decomposed into four Gaussians, three for the host line profile and one for outflow (green). The black line includes all four Gaussian components. The insert compares the sum of the outflow fits to an  $r = 0.6''$  aperture spectrum. *Center left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. *Center and bottom right:* position-velocity diagrams for two orthogonal position angles. Molecular outflow is seen strongly on the red and weakly on the blue side of the line core, with spatial shifts better matching the PA +45° PV diagram.



**Fig. A.13.** NOEMA CO(1-0) data for NGC 7479. *Top left:* moment-zero map. *Top right:* spectrum in a  $r = 3''$  aperture centered on the continuum nucleus. The spectrum is fitted with three Gaussians for the host line profile. There is no outflow detection. *Bottom left:* center positions and their errors, from UV fitting a Gaussian model to individual velocity channels. The dotted line indicates the N-S direction of NGC 7479's bar. *Bottom right:* position-velocity diagram. The pronounced bipolar structure in the PV diagram is ascribed to flow in the galaxy's bar, and we adopt an upper limit for molecular outflow.

**Appendix B: Detected lines**

ALMA bands. Table B.1 lists these lines and their measured total fluxes.

The observations used in this paper were targeted at CO(1-0) but some additional lines are detected in the wide NOEMA or

**Table B.1.** Detected lines.

Source	$\nu_{\text{Obs}}$ (GHz)	Line	$\nu_{\text{Rest}}$ (GHz)	Flux (Jy km s <sup>-1</sup> )
III Zw 035	112.21	CO(1-0)	115.27	70.6
III Zw 035	110.45	CN N=1-0	113.49 m	5.8
III Zw 035	110.15	CN N=1-0	113.19 m	2.7
IRAS F05189–2524	110.54	CO(1-0)	115.27	25.7
IRAS F05189–2524	108.83	CN N=1-0	113.49 m	4.4
IRAS F05189–2524	108.52	CN N=1-0	113.19 m	2.3
IRAS F05189–2524	95.96	HC <sub>3</sub> N 11–10	100.08	0.5
NGC 2623	113.18	CO(1-0)	115.27	161.0
NGC 2623	111.44	CN N=1-0	113.49 m	8.5
NGC 2623	111.13	CN N=1-0	113.19 m	4.9
IRAS F09111–1007 W	109.33	CO(1-0)	115.27	51.0
IRAS F09111–1007 W	107.64	CN N=1-0	113.49 m	4.0
IRAS F09111–1007 W	107.33	CN N=1-0	113.19 m	2.6
IRAS F09111–1107 W	94.92	HC <sub>3</sub> N 11–10	100.08	0.3
IRAS F09111–1007 E	109.28	CO(1-0)	115.27	19.7
IRAS F10173+0828	109.89	CO(1-0)	115.27	24.2
IRAS F10173+0828	108.19	CN N=1-0	113.49 m	1.4
IRAS F10173+0828	107.89	CN N=1-0	113.19 m	0.7
IRAS F10173+0828	95.41	HC <sub>3</sub> N 11–10	100.08	0.7
IRAS F12224–0624	112.30	CO(1-0)	115.27	22.6
IRAS F12224–0624	110.56	CN N=1-0	113.49 m	2.0
IRAS F12224–0624	110.25	CN N=1-0	113.19 m	1.1
IRAS F12224–0624	98.02	?	100.62 m	0.4
IRAS F12224–0624	97.50	HC <sub>3</sub> N 11–10	100.08	0.8
NGC 4418	114.47	CO(1-0)	115.27	117.0
NGC 4418	112.70	CN N=1-0	113.49	10.5
NGC 4418	112.38	CN N=1-0	113.19	4.6
NGC 4418	101.81	CH <sub>3</sub> CCH	102.54	0.7
IC 860	113.80	CO(1-0)	115.27	73.1
IC 860	112.05	CN N=1-0	113.49 m	4.6
IRAS 13120–5453	111.80	CO(1-0)	115.27	234.0
IRAS 13120–5453	110.07	CN N=1-0	113.49 m	24.6
IRAS 13120–5453	109.75	CN N=1-0	113.19 m	14.6
IRAS 13120–5453	97.58	?	100.61 m	0.8
IRAS 13120–5453	97.08	HC <sub>3</sub> N 11–10	100.08	1.0
IRAS F14378–3651	107.92	CO(1-0)	115.27	24.7
IRAS F14378–3651	106.26	CN N=1-0	113.49 m	2.4
IRAS F14378–3651	105.95	CN N=1-0	113.19 m	1.3
IRAS F17207–0014	110.54	CO(1-0)	115.27	171.0
IRAS F17207–0014	108.83	CN N=1-0	113.49 m	10.0
IRAS F17207–0014	108.53	CN N=1-0	113.19 m	4.5
IRAS F17207–0014	98.33	CH <sub>3</sub> CCH	102.54	0.2
IRAS F17207–0014	96.50	?	100.63 m	0.5
IRAS F17207–0014	95.97	HC <sub>3</sub> N 11–10	100.08	2.9
IRAS F20551–4250	110.52	CO(1-0)	115.27	54.9
IRAS F20551–4250	108.82	CN N=1-0	113.49	2.3
IRAS F20551–4250	108.50	CN N=1-0	113.19	1.0
NGC 7479 (30'')	114.39	CO(1-0)	115.27	270.0
NGC 7479 (30'')	112.60	CN N=1-0	113.49 m	6.4

**Notes.** Total fluxes, or fluxes in the indicated aperture diameter, for lines detected in the spectral ranges covered by NOEMA and ALMA. For the very rich spectrum of NGC 4418 we list only a few bright lines that are also detected in other targets. The listed frequencies are indicative for identification purposes only, in particular for lines with multiple components marked “m”.

## Appendix C: Additional table

Table C.1. Other properties of combined sample.

Source	$D_L$ (Mpc)	$\log(\dot{M}_{\text{out}})$ ( $M_\odot$ )	$t_{\text{flow}}$ (Myr)	$\dot{M}_{\text{out}}$ ( $M_\odot \text{ yr}^{-1}$ )	$\log(P_{\text{kin}})$ ( $\text{erg s}^{-1}$ )	$\log(\dot{M}_{\text{out}} v)$ ( $\text{g cm s}^{-2}$ )	$\log(\Sigma_{\text{FIR}})$ ( $L_\odot \text{ kpc}^{-2}$ )	$\log(L_{\text{Bol}})$ ( $L_\odot$ )	$\log(L_{\text{AGN}})$ ( $L_\odot$ )	Ref. (AGN)
NGC 253	3.40	7.20	4.30	3.71	39.47	33.07	12.14	10.58	<7.07	Cicone et al. (2014)
IZw 1	273.98	<7.62					10.87	12.01	11.96	Veilleux et al. (2013)
III Zw 035	119.91	8.59	6.88	56.70	42.24	35.05	11.81	11.71	<7.76	González-Martín et al. (2015)
PG 0157+001	781.80	7.99	1.54	63.81	42.51	35.21	>11.18	12.68	12.50	Veilleux et al. (2013)
NGC 1068	14.40	6.93	1.02	8.27	40.42	33.72	11.36	11.37	10.15	Ricci et al. (2017)
NGC 1266	29.90	7.67	1.21	38.31	42.10	34.89	>11.89	10.54	9.73	Cicone et al. (2014)
IRAS F03158+4227	632.90	8.88	6.92	110.83	43.43	35.79		12.68	12.36	Fluetsch et al. (2019)
NGC 1377	21.00	6.88	1.42	5.36	40.51	33.67	>11.49	10.09	9.35	Cicone et al. (2014)
NGC 1433	9.90	5.82	0.94	0.70	39.34	32.64		9.24	8.32	Fluetsch et al. (2019)
NGC 1614	68.88	7.51	1.52	21.33	41.94	34.68	11.21	11.73	<8.94	Brightman & Nandra (2011)
IRAS F05081+7936	239.20	7.98	1.22	28.23	43.20	35.60	10.61	12.05	<8.79	Fluetsch et al. (2019)
IRAS F05189-2524	189.47	7.87	2.30	32.56	42.31	34.96	11.78	12.22	12.09	Veilleux et al. (2013)
NGC 2146	17.20	7.94	10.20	8.49	41.03	34.03	10.46	11.16	<7.51	Cicone et al. (2014)
NGC 2623	80.03	7.05	1.01	11.21	42.09	34.62	11.93	11.62	10.62	Evans et al. (2008)
NGC 2623 fossil		6.66	12.80	0.36	40.26	32.96				
IRAS F08572+3915	260.32	8.56	1.00	363.80	43.87	36.26	11.73	12.20	12.06	Veilleux et al. (2013)
IRAS F09111-1007 W	242.11	8.01	1.81	57.12	42.29	35.08	11.64	12.01	<10.80	Nardini et al. (2010)
M 82	3.90	7.86	9.78	7.34	40.36	33.67	11.08	10.90	<7.96	Cicone et al. (2014)
IRAS F10035+4852	291.20	<8.45						12.06	<11.52	Fluetsch et al. (2019)
IRAS F10173+0828	217.64	<7.29					11.77	11.90	<8.18	Iwasawa et al. (2011)
NGC 3256 N	35.00	7.26	0.57	31.93	42.69	35.15	10.79	11.60	<8.39	Cicone et al. (2014)
NGC 3256 S		7.13	1.33	10.11	41.90	34.50				
IRAS F10565+2448	190.70	8.34	2.39	91.90	42.77	35.42	11.54	12.12	11.79	Veilleux et al. (2013)
IRAS F11119+3257	920.20	8.75	6.85	81.43	43.41	35.71		12.72	12.63	Veilleux et al. (2013)
NGC 3628	7.70	6.98	4.46	2.13	39.74	33.08	11.13	10.08	<7.21	Cicone et al. (2014)
ESO 320-G030	46.60	6.81	2.36	2.76	41.00	33.77	11.28	11.33	<9.33	Alonso-Herrero et al. (2012)
IRAS F12112+0305 SW	330.34	7.53	2.31	14.62	41.68	34.47		12.38	11.64	Veilleux et al. (2013)
IRAS F12112+0305 NE		8.80	3.36	188.41	43.11	35.74				
IRAS F12224-0624	115.71	6.95	0.46	19.42	42.29	34.84	11.83	11.35	<10.75	This work
NGC 4418	25.35	5.73	0.09	6.29	41.32	34.11	12.50	10.94	10.68	Veilleux et al. (2013)
Mrk 231	186.42	8.43	0.84	323.84	43.70	36.15	>12.26	12.61	12.51	Veilleux et al. (2013)
IC 860	56.11	<6.55					12.07	11.19	<9.19	Alonso-Herrero et al. (2012)
IRAS 13120-5453	136.14	8.18	1.35	111.53	42.50	35.33	11.99	12.33	11.85	Veilleux et al. (2013)
Mrk 273	166.43	8.21	0.83	196.30	43.42	35.91	>12.13	12.24	11.77	Veilleux et al. (2013)
4C 12.50	568.50	<8.11						12.41	12.11	Fluetsch et al. (2019)
SDSS J135646.10+102609.0	574.70	8.09	0.59	207.49	43.21	35.82		12.43	12.41	Sun et al. (2014)
IRAS F14348-1447 SW	375.26	8.76	3.80	151.34	42.92	35.60		12.42	11.66	Veilleux et al. (2013)
IRAS F14348-1447 NE		8.05	2.44	45.91	42.30	35.03				
IRAS F14378-3651	306.70	<7.87					11.81	12.26	11.59	Veilleux et al. (2013)
Zw 049.057	56.40	6.96	0.26	35.09	42.00	34.82	11.90	11.26	<9.26	Alonso-Herrero et al. (2012)
Arp 220	78.70	7.02	0.19	54.93	42.40	35.12	12.53	12.24	11.01	Veilleux et al. (2013)
Mrk 876	605.21	9.38	3.31	720.05	44.40	36.68	10.70	12.30	12.17	Ricci et al. (2017)
NGC 6240	106.81	8.59	1.43	272.00	43.24	35.89	11.28	11.91	11.66	Veilleux et al. (2013)
IRAS F17020+4544	270.52	8.27	0.59	315.95	44.44	36.52		11.61	11.13	Gioletti et al. (2017)
IRAS F17132+5313	226.60	<8.16						12.00	8.70	Iwasawa et al. (2011)
IRAS F17207-0014	189.20	7.52	3.33	98.96	43.59	35.84	11.97	12.48	<11.18	Veilleux et al. (2013)
PDS 456	898.60	8.61	3.37	120.03	43.39	35.83		13.42	13.41	Bischetti et al. (2019)
NGC 6764	32.00	6.44	20.00	0.14	38.09	31.67		10.46	8.65	Cicone et al. (2014)
IRAS F20100-4156	608.20	8.86	2.43	295.08	43.90	36.24		12.73	12.16	Spoon et al. (2013)
IC 5063	48.71	7.23	0.90	18.61	42.17	34.77	10.24	10.92	10.84	Ricci et al. (2017)
IRAS F20551-4250	190.20	7.43	0.46	59.44	42.65	35.26	11.64	12.11	11.87	Veilleux et al. (2013)
IRAS F22491-1808	352.74	8.13	1.55	86.73	42.46	35.25	11.26	12.26	11.42	Veilleux et al. (2013)
NGC 7479	34.31	<6.22					11.31	10.89	9.95	Ricci et al. (2017)
IRAS F23060+0505	834.25	9.42	5.66	463.76	43.85	36.31		12.59	12.48	Cicone et al. (2014)
IRAS 23365+3604	289.17	8.12	2.67	49.46	42.50	35.15	11.65	12.27	11.92	Veilleux et al. (2013)

**Notes.** Molecular outflow masses, flow times, and outflow rates for the combined sample, as computed using the observed quantities from Tables 3 and 4. We also list adopted distances, far-infrared surface brightnesses (from Lutz et al. 2016, 2018), and bolometric and AGN luminosities. Some references quote AGN fractions to the bolometric luminosity which we have converted to  $L_{\text{AGN}}$  for consistency. The limit for IRAS F12224-0624 was estimated from *Spitzer*-IRS data with methods as used by Veilleux et al. (2013). We note that AGN luminosities are a mixture of X-ray-based (reacting rapidly to accretion variations) and mid-infrared-based (averaging over typically tens of years).